

# **Centre-to-Limb Variation Calculations with 3D Model Stellar Atmospheres**

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# 1D model atmospheres

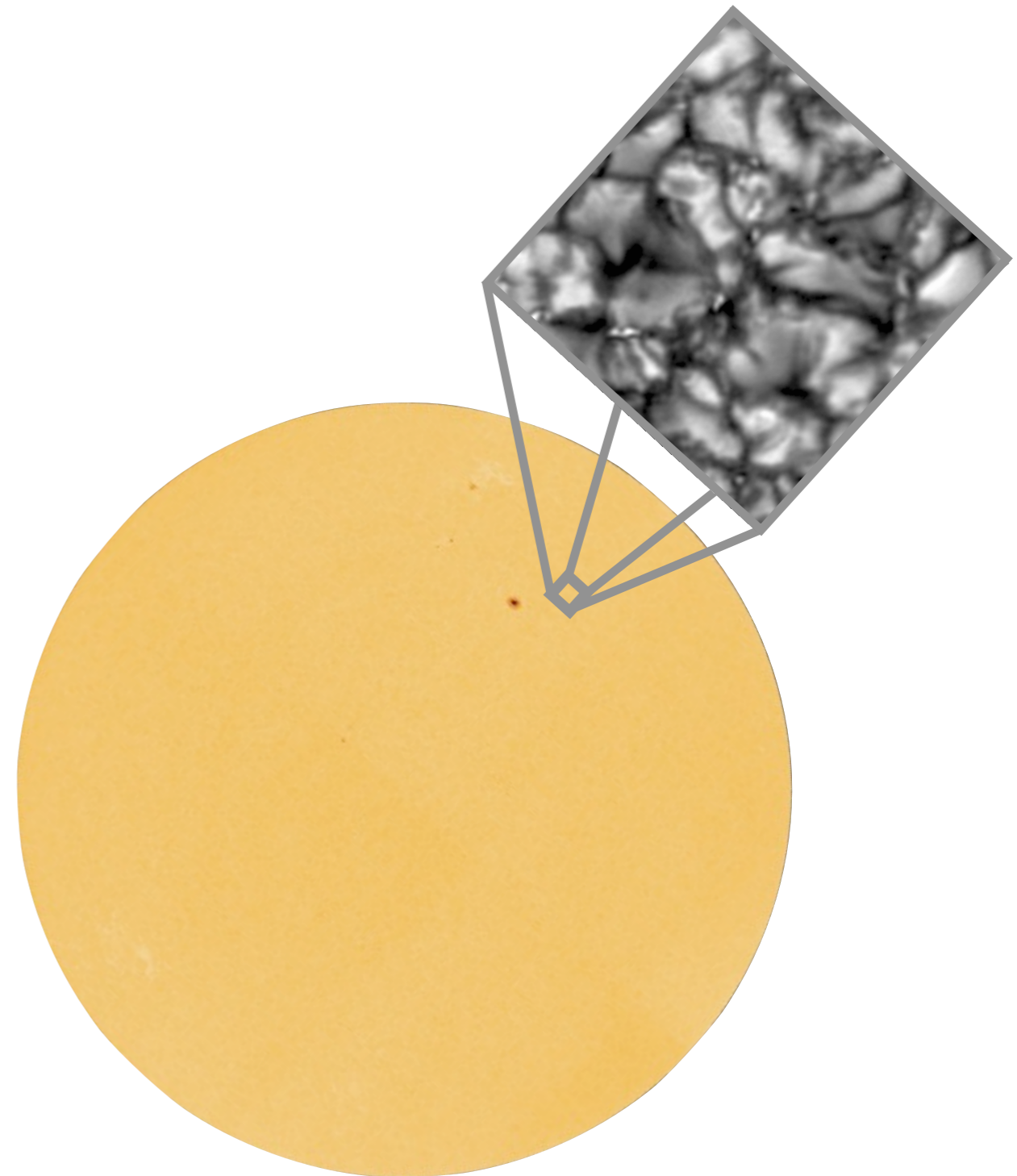
- Classical model atmospheres of FGKM stars (MARCS, Gustafsson et al. 2008; ATLAS, Kurucz 1979; NextGen/PHOENIX, Hauschildt et al. 1999)
- 1D Homogeneous stratification
- Hydrostatic
- Stationary
- Radiative transfer: 100,000 wavelengths or more
- Convection: mixing-length theory, free parameters

# 3D models

- **STAGGER** (Nordlund & Galsgaard 1995; Collet et al. 2018); CO5BOLD (Freytag et al. 2012); MURaM (Vögler et al. 2005)
- Solution of mass, momentum, and energy conservation equations
- 3D geometry
- Time-dependent
- 3D non-grey radiative transfer (with multi-group opacities)
- (Magnetic fields)
- Convection: no need for dedicated free parameters

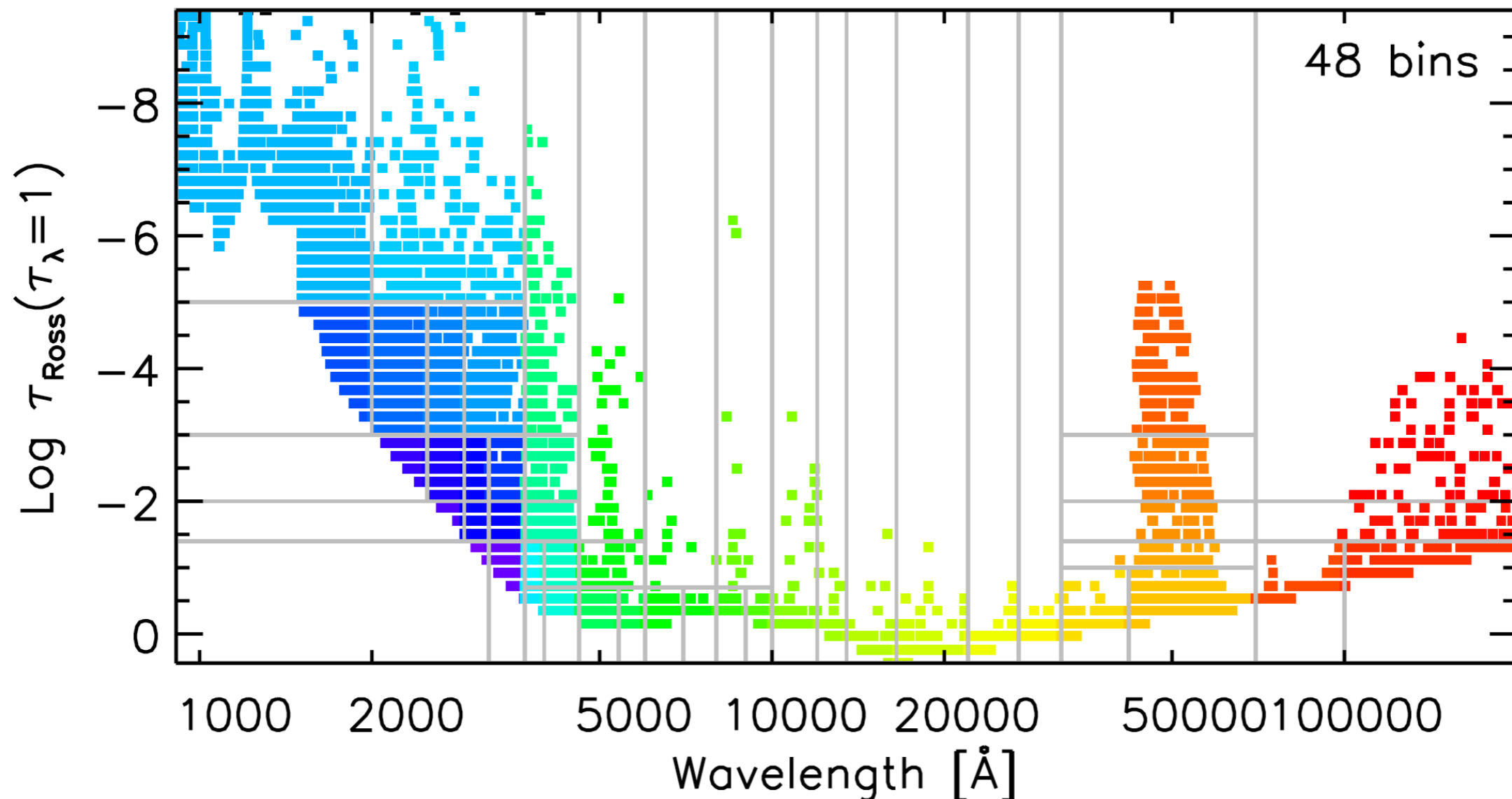
# Simulations: setup and input physics

- “Box in the star”
- Equation of state: Mihalas et al. 1988 with updates
- Opacities: Uppsala/MARCS package (Gustafsson et al. 1975, 2008; B. Plez et al.)
- Radiative transfer:  $\sim 10$ - $20$  inclined rays per surface point

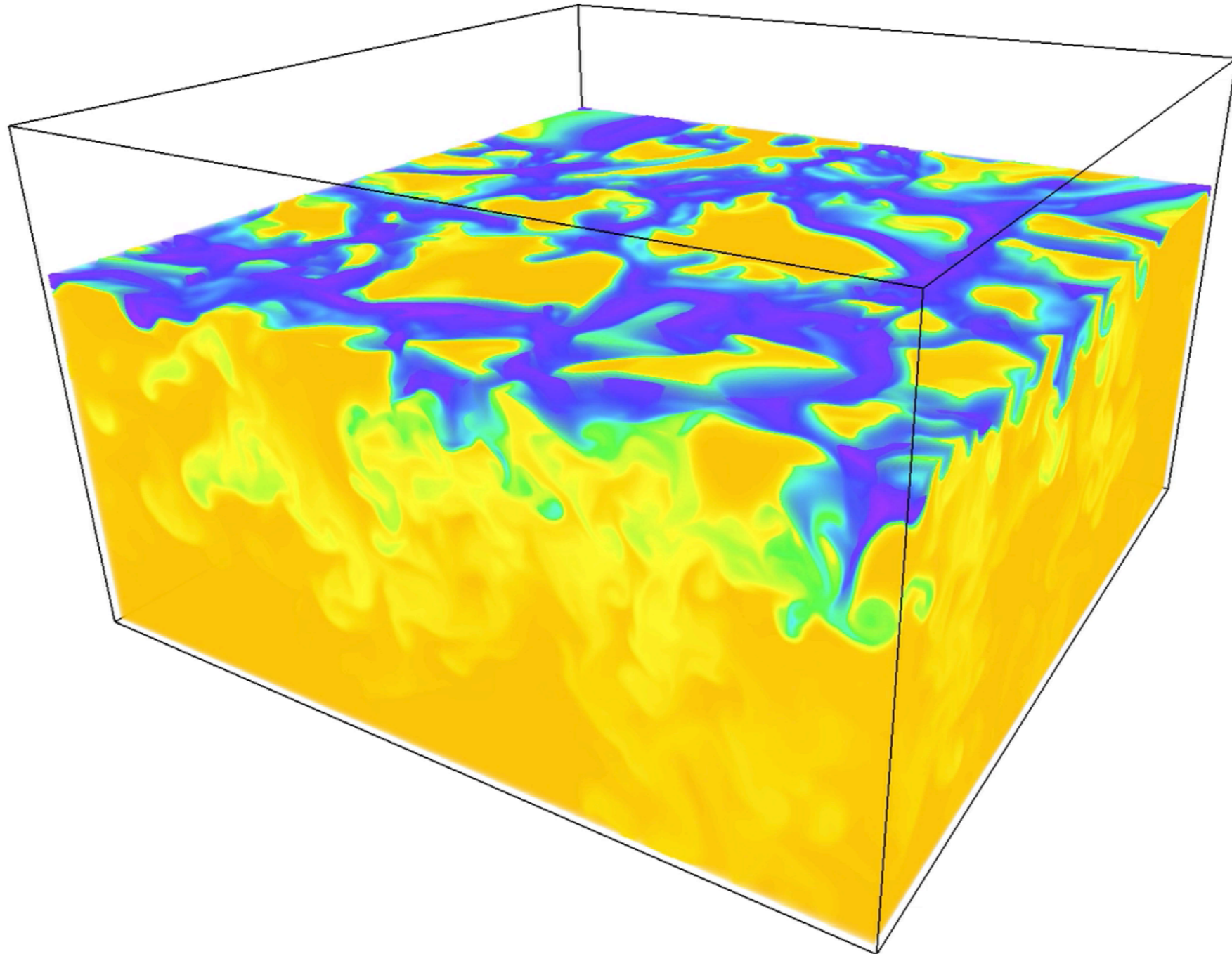


# Opacity binning

- Sort monochromatic wavelengths into groups (opacity bins)
- Solve radiative transfer for average opacities and integrated source functions in bins

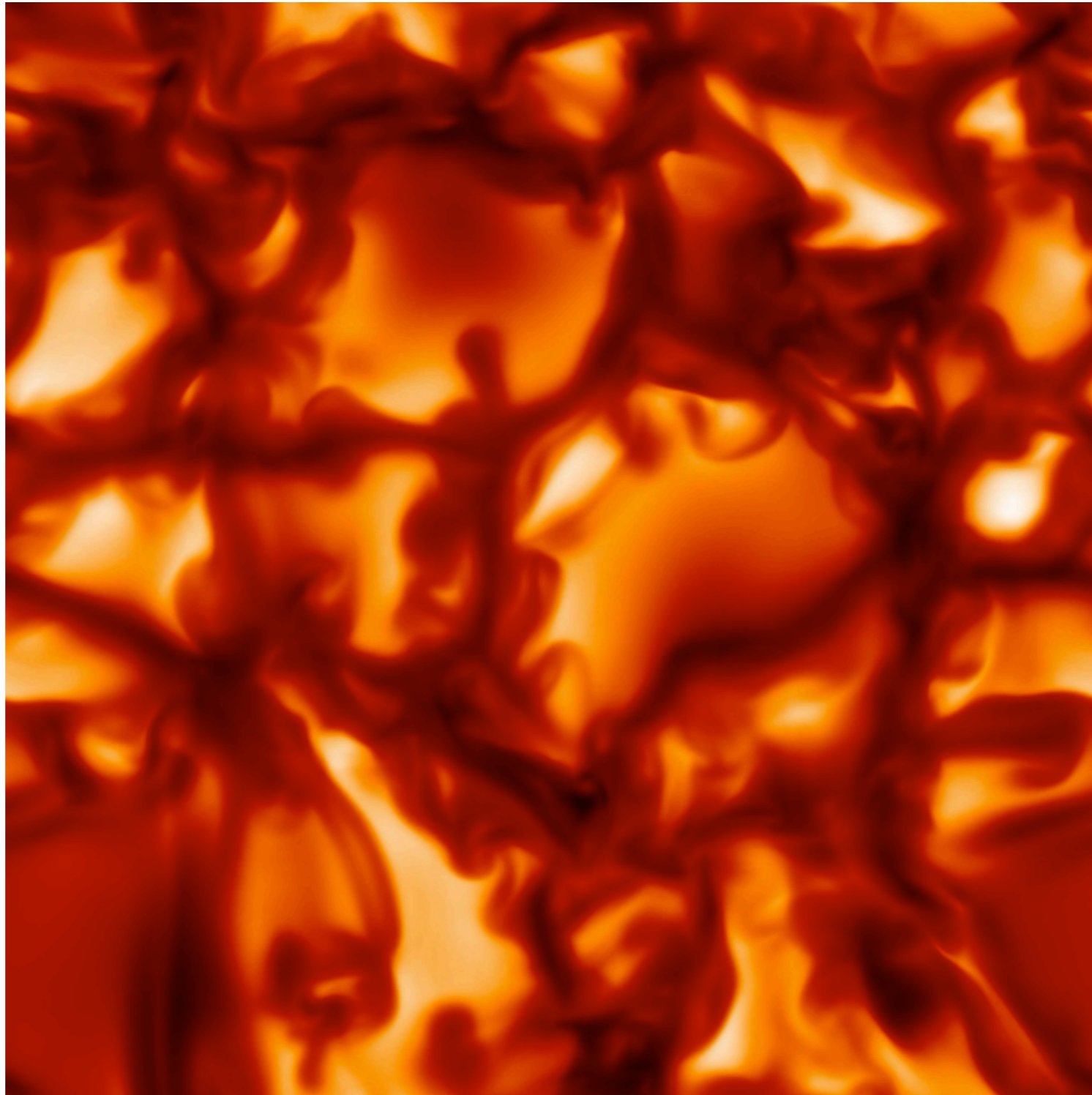


# 3D simulations: surface convection



**STAGGER solar surface convection simulation (R. Collet)**

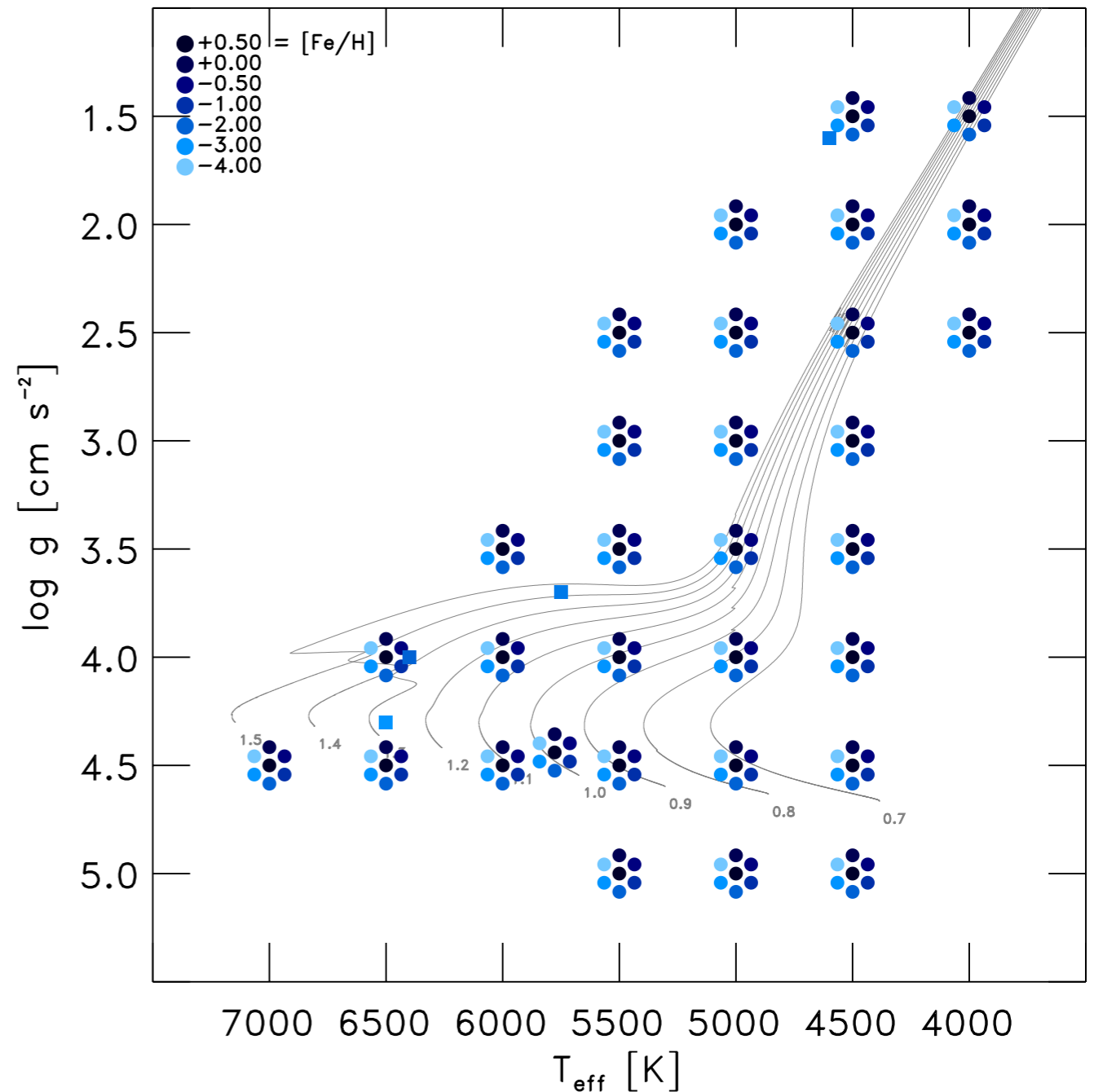
# Surface intensity



**Solar simulation, 6 Mm x 6 Mm, (R. Collet)**

# Grid of 3D simulations

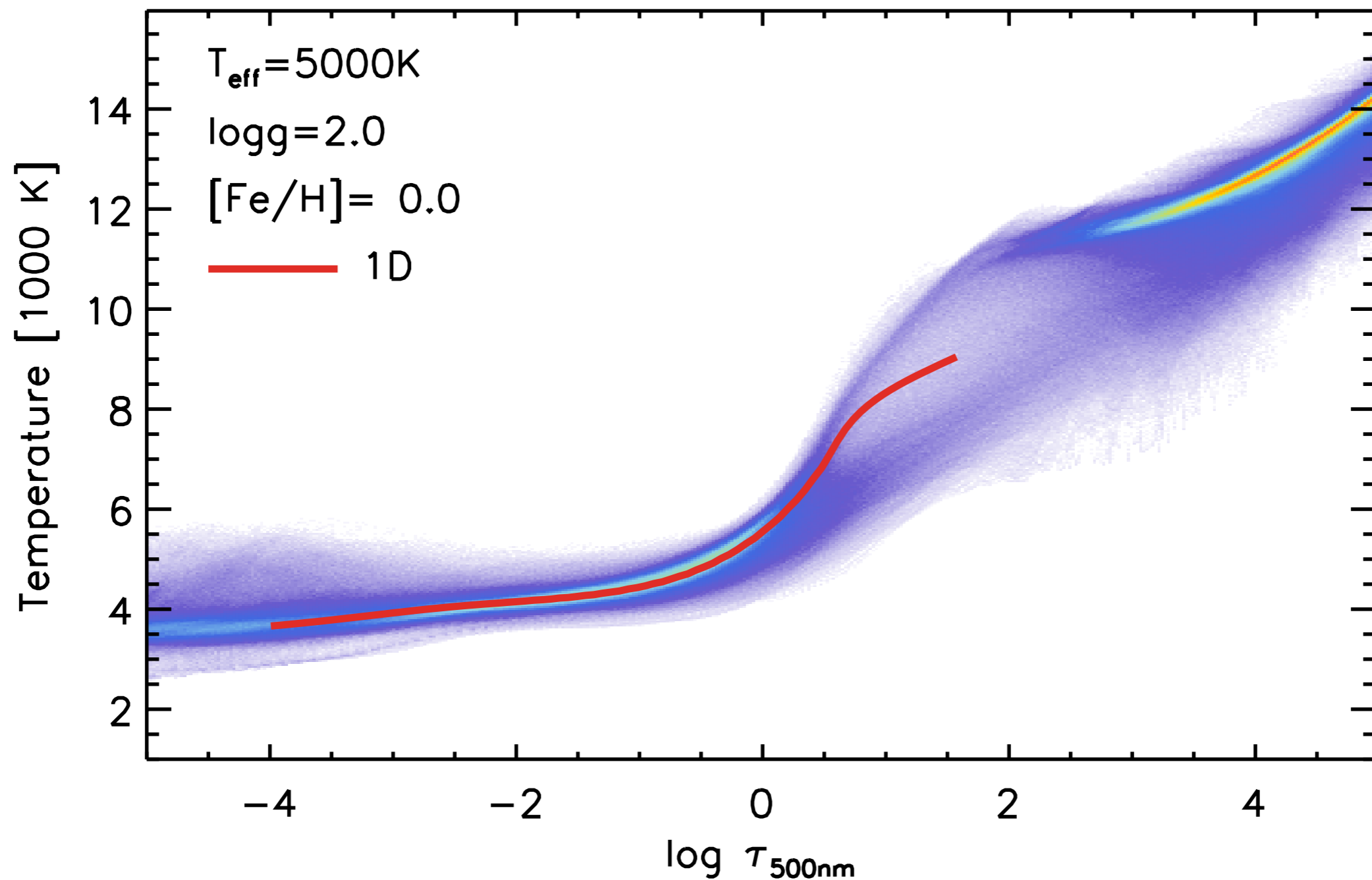
- STAGGER (Collet et al. 2011; Magic et al. 2013;): ~200 3D convection simulations of FGK stars
- Trampedach et al. (2013); CO<sup>5</sup>BOLD/CIFIST (Tremblay et al. 2013); MURaM (Beeck et al. 2013)
- Systematic study of 3D-1D differences in synthetic spectra and colours





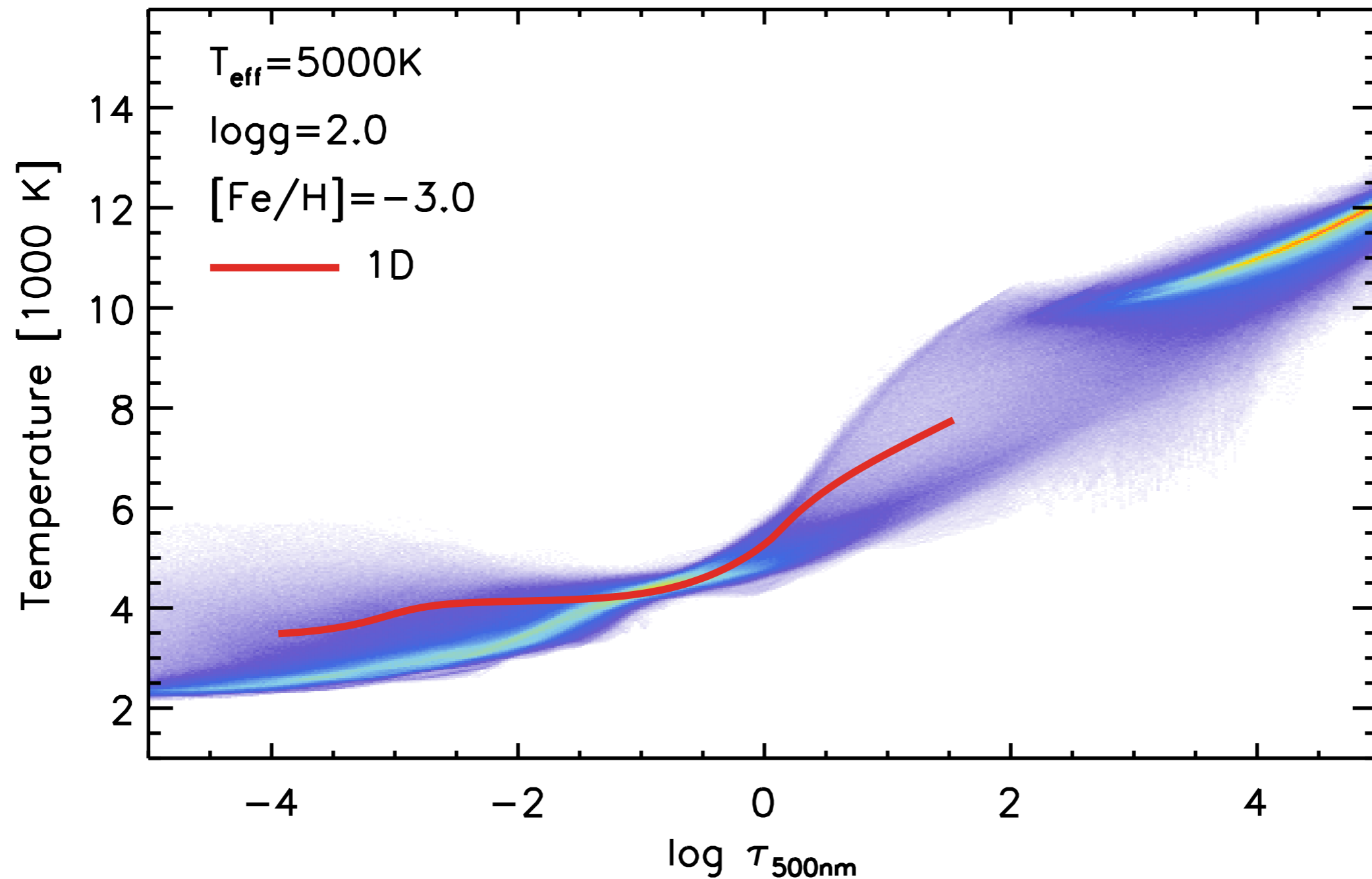
# Solar metallicity

## Temperature vs optical depth



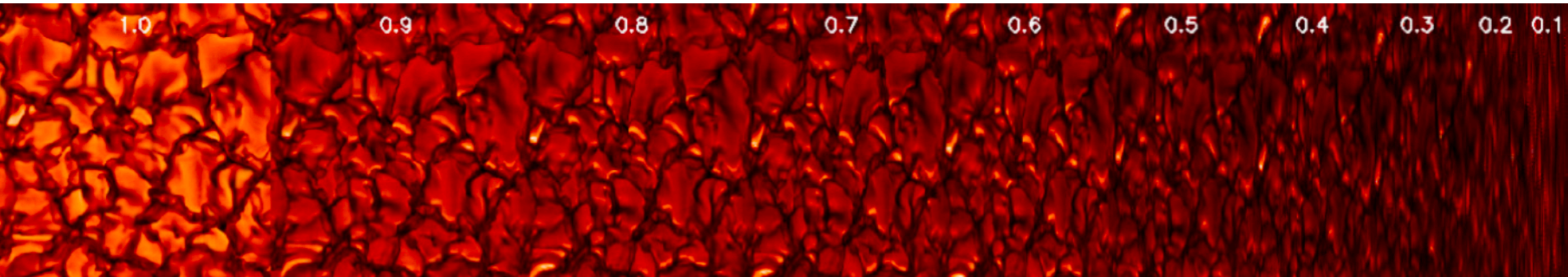
# Low metallicity

## Temperature vs optical depth



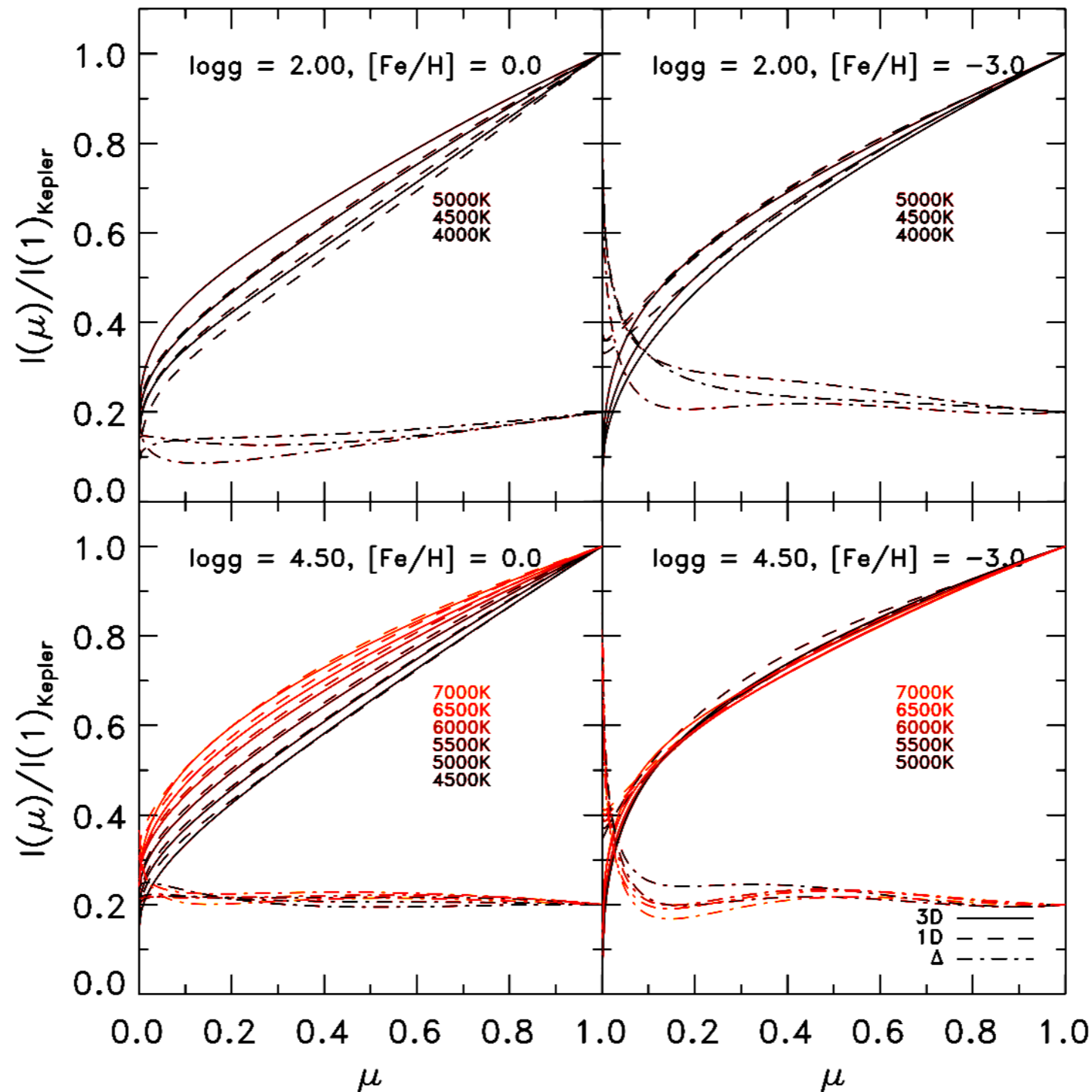
# Limb darkening

- The diameter of CoRoT target HD49933: limb-darkening calculations with 3D models combined with interferometric and asteroseismic observations (Bigot et al. 2011)



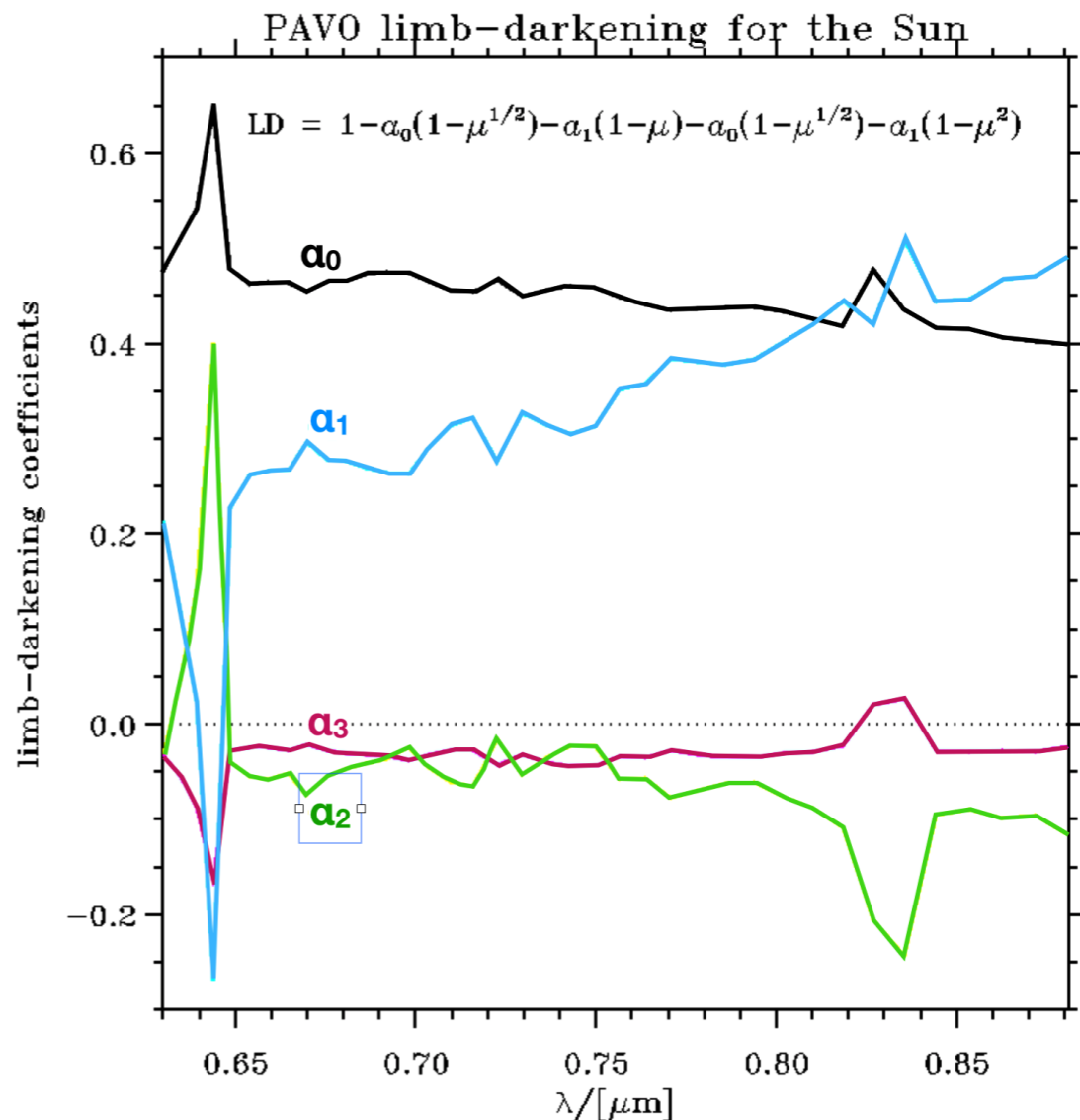
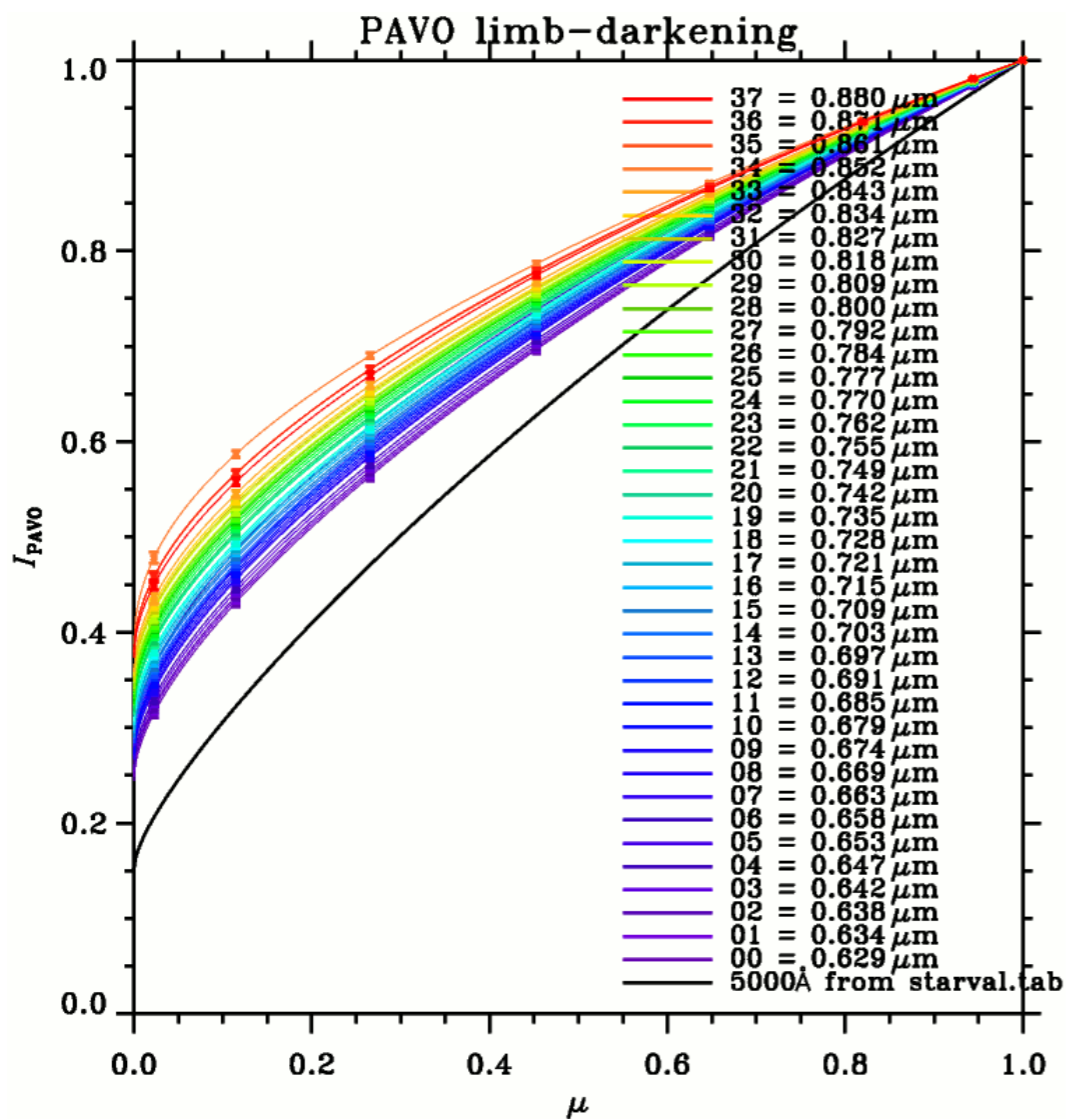
# Limb darkening, 1D-3D

Centre-to-limb variations, Kepler band, 3D vs 1D (STAGGER; Magic et al. 2015)

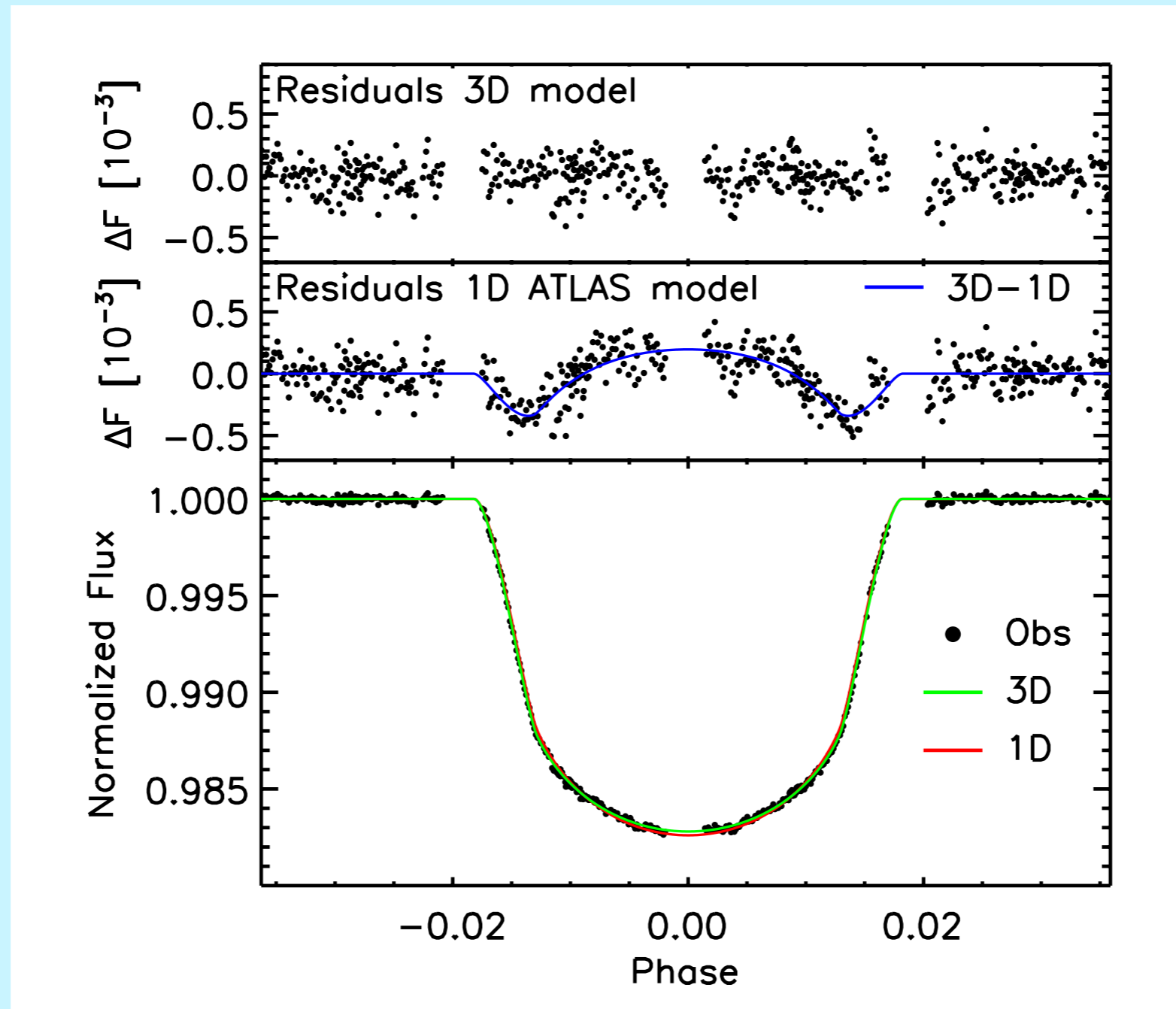


# Limb-darkening coefficients

- 3D solar simulation: fit to 4-term Claret (2000) limb-darkening law for 38 CHARA/PAVO channels (Trampedach et al. in prep.)
- Wavelength dependence (especially  $\alpha_1$  coefficient)
- Strong molecular features at 0.64  $\mu\text{m}$  and at 0.83  $\mu\text{m}$



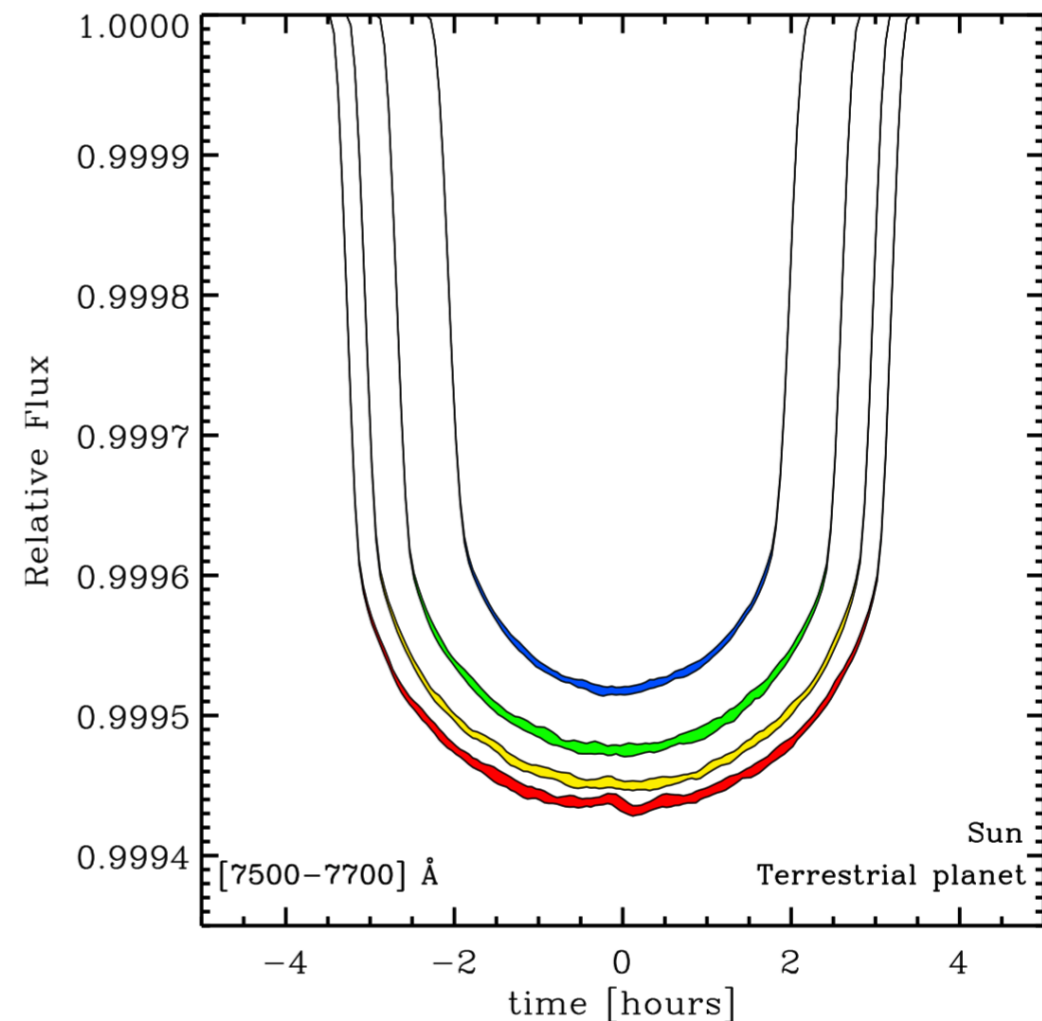
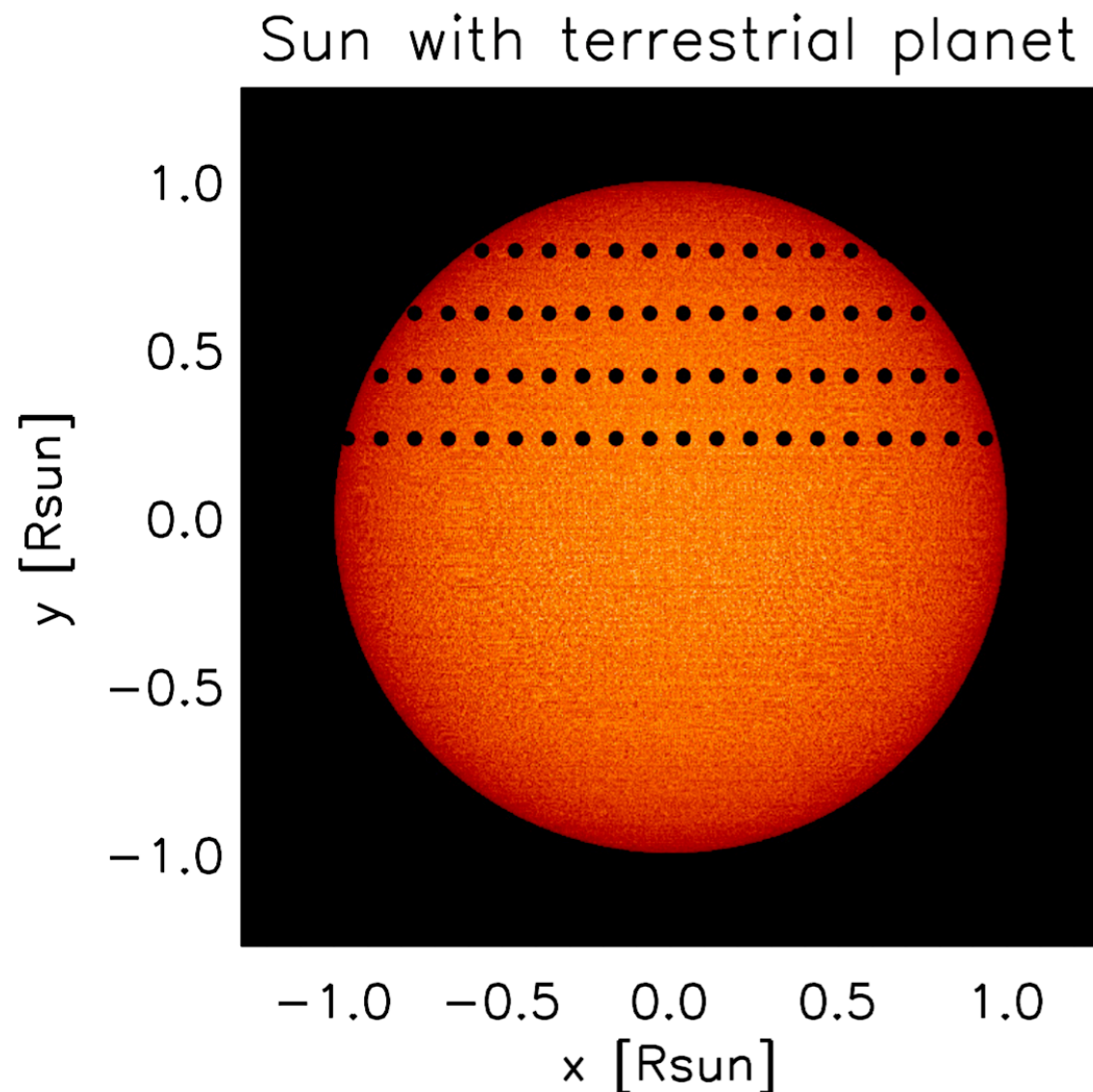
# Exoplanet transits



HD209458b transit, HST 2900 Å - 5700 Å band (Hayek et al. 2012)

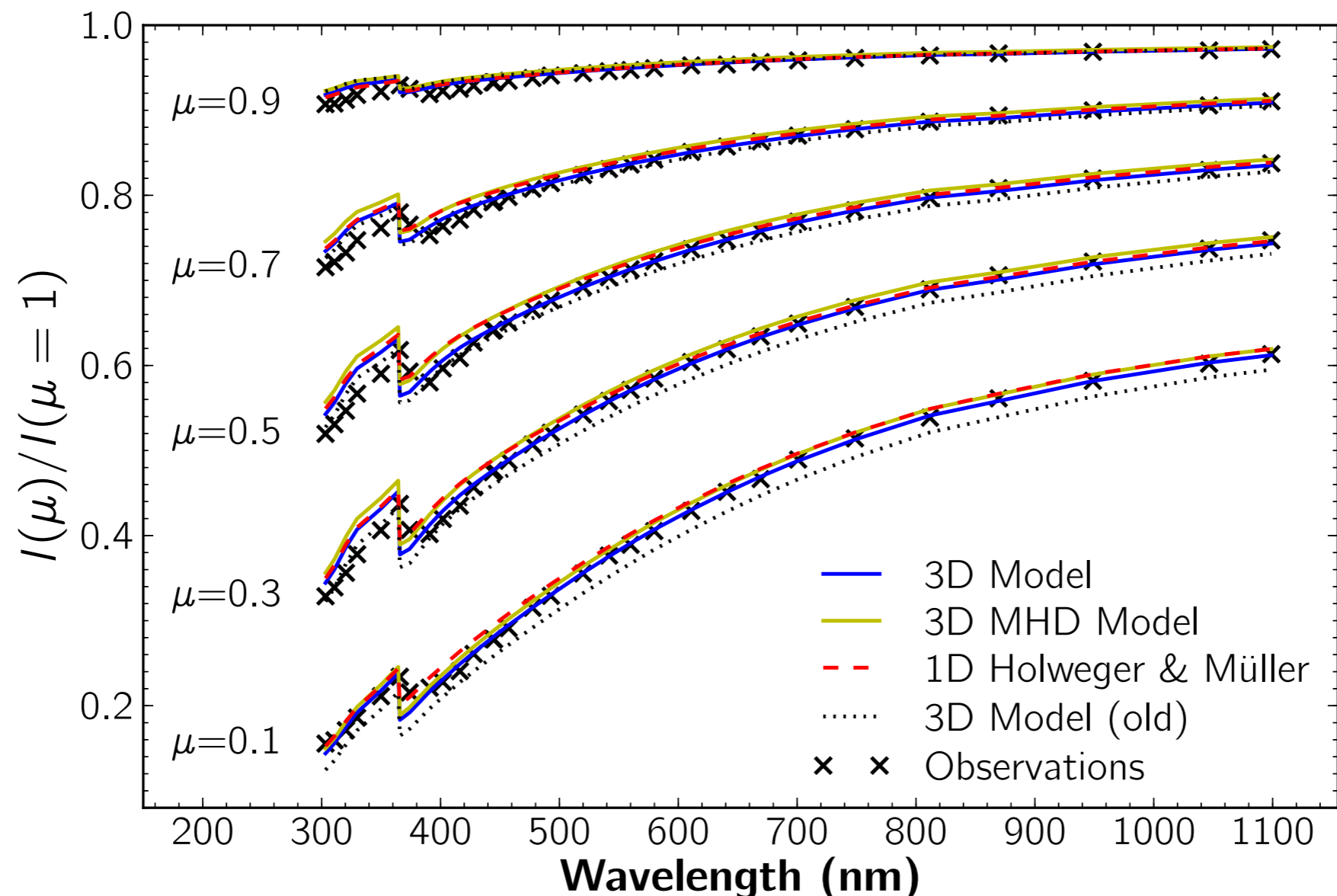
# Exoplanet transits (2)

- Simulated transits with synthetic stellar disc images from 3D simulations: granulation noise



# Limb darkening: Sun

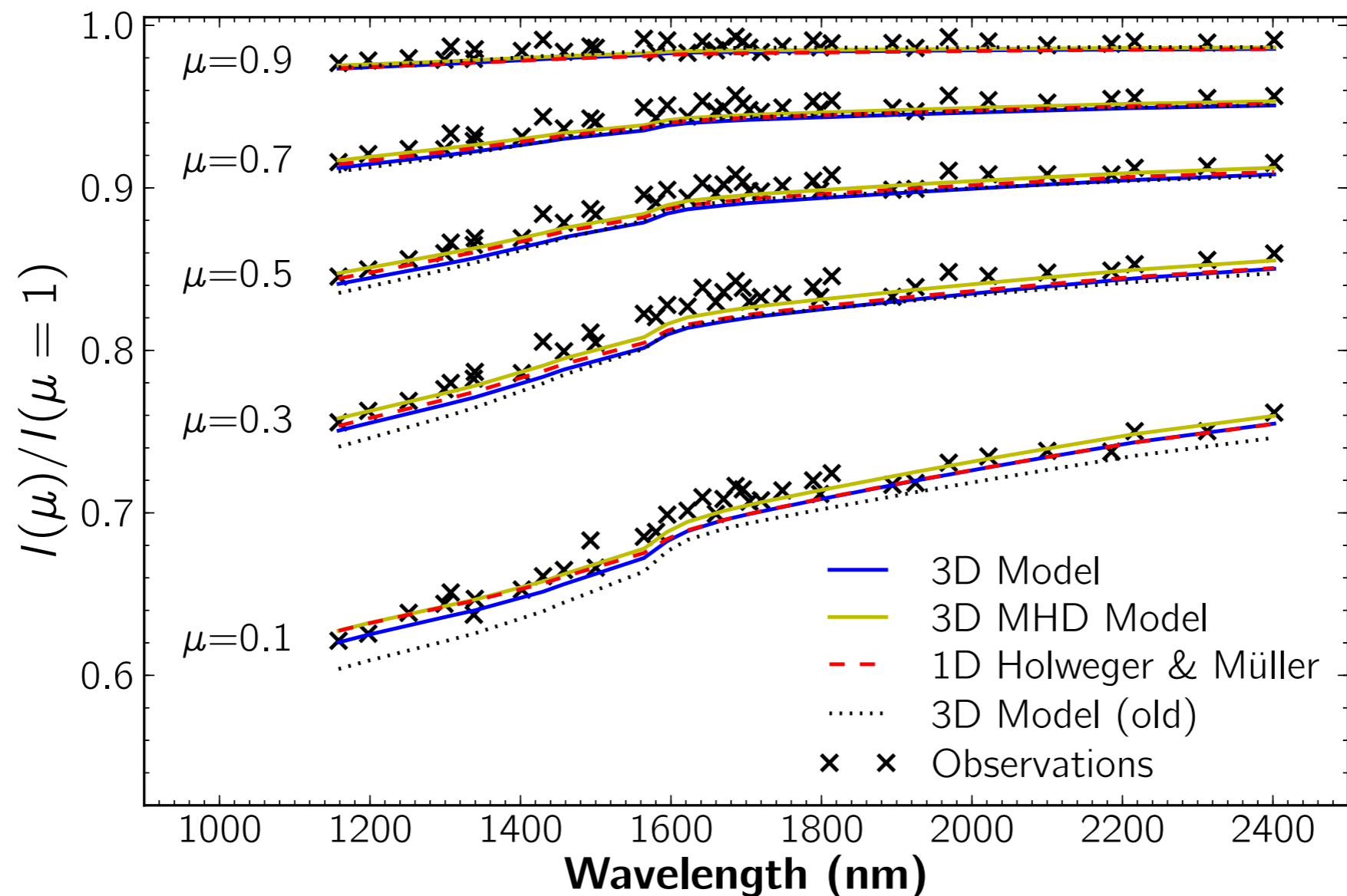
- Very good agreement between limb-darkening curves from 3D models and observations at UV and visible wavelengths (Pereira, Asplund, Collet et al. 2013; data: Pierce & Slaughter 1977; Neckel & Labs 1994)





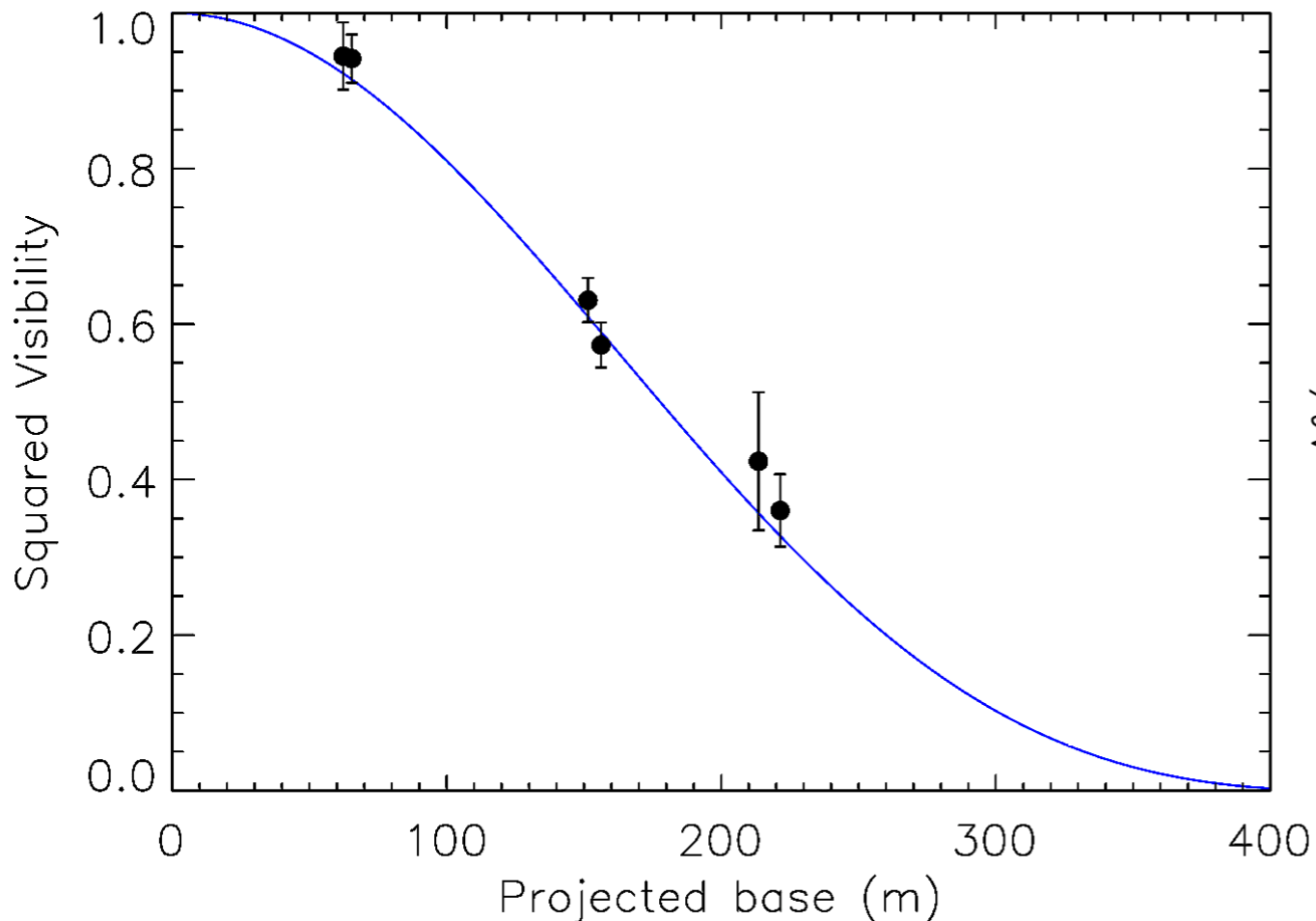
# Limb darkening: Sun (2)

- Poorer agreement with observations at long wavelengths (Pereira et al. 2013; data: Pierce et al. 1977)

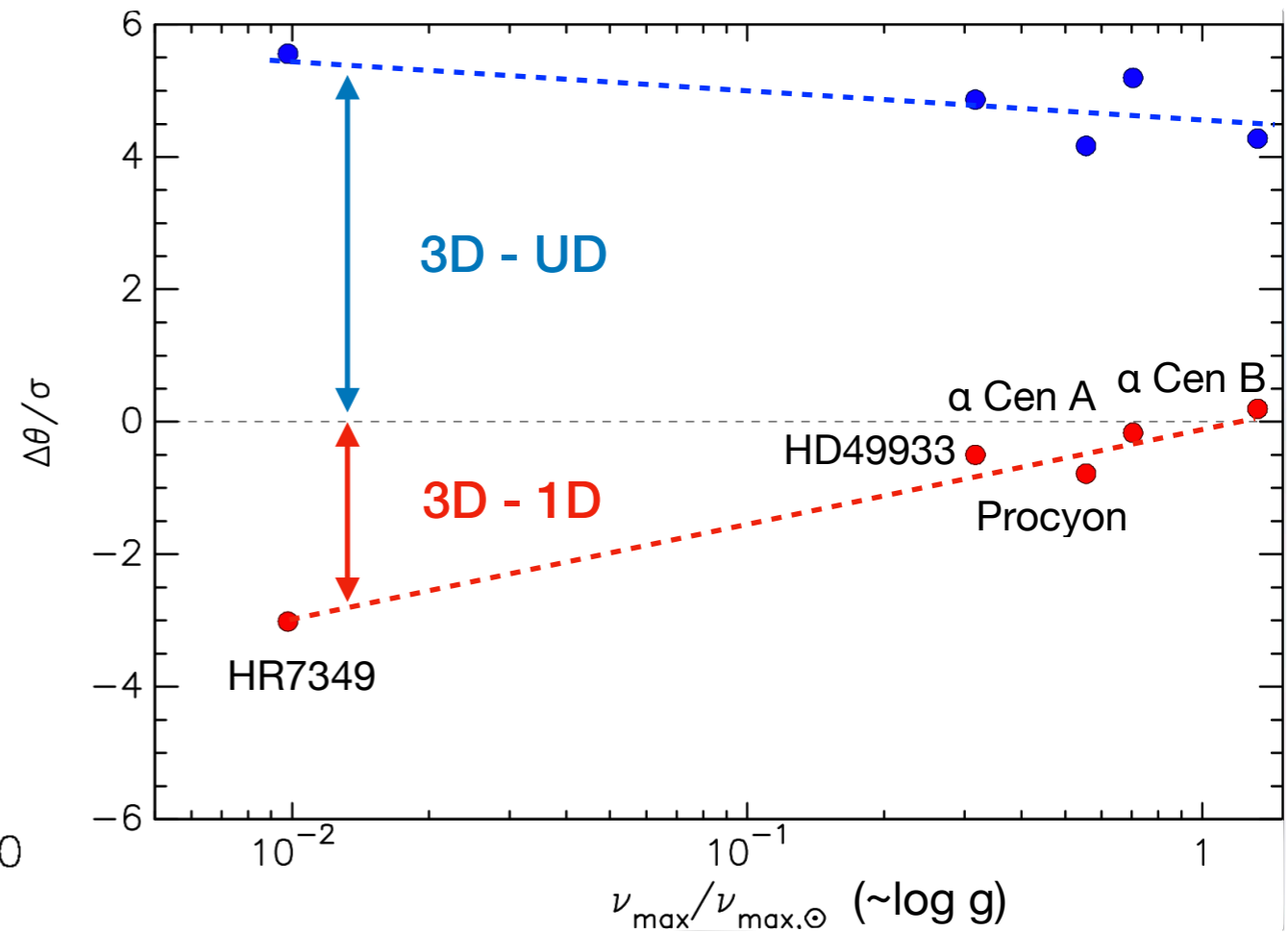


# Interferometric stellar angular diameters

- In general:  $\theta_{UD} < \theta_{3D} < \theta_{1D}$
- Significant ( $3\sigma$ ) 3D-1D limb-darkening corrections for giants



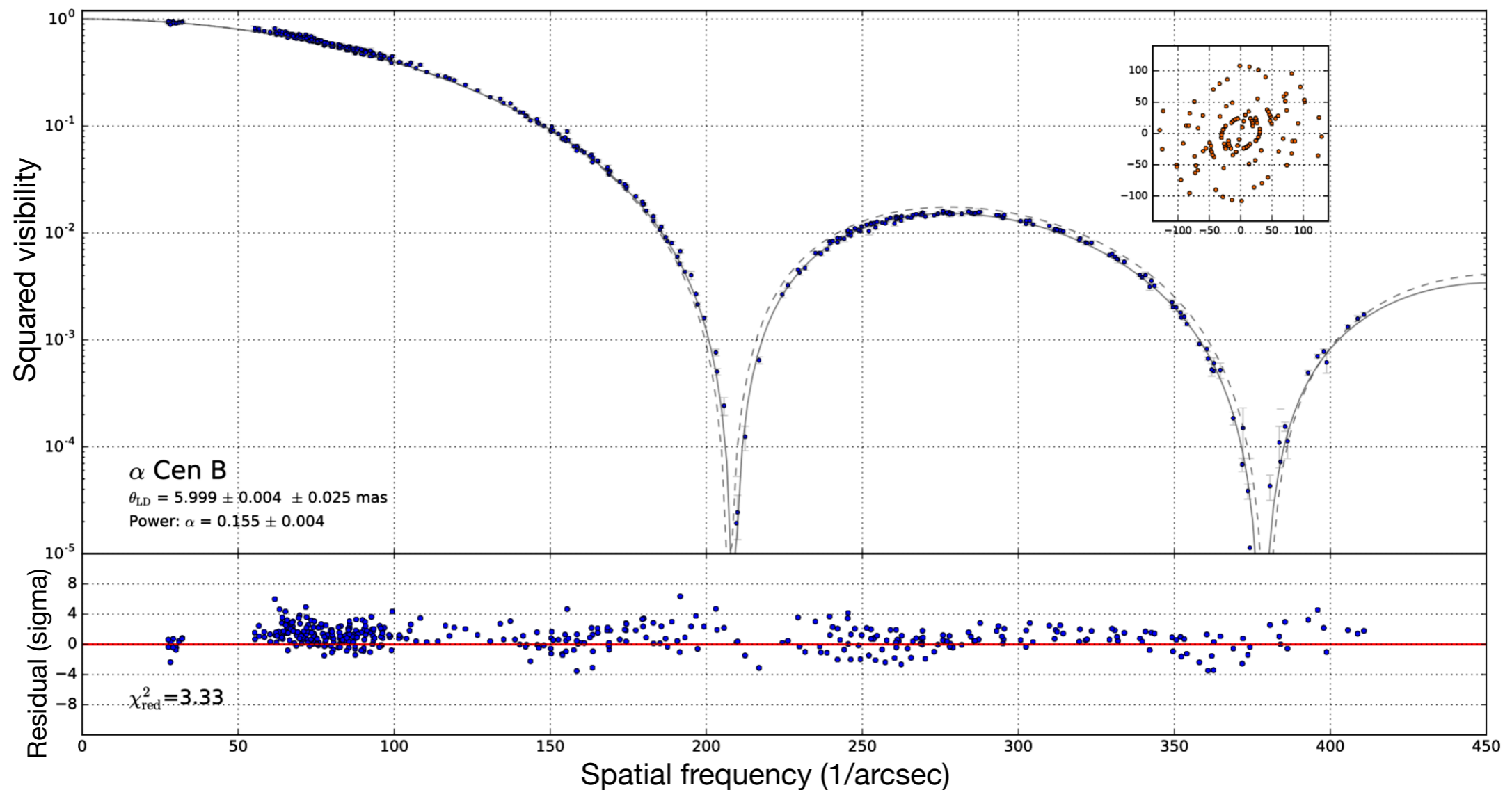
HD49933 (Bigot et al. 2011)



Bigot et al., in prep; Chiavassa et al. 2012;  
Bigot et al. 2011; Kervella et al. 2017; Bigot et al. 2006

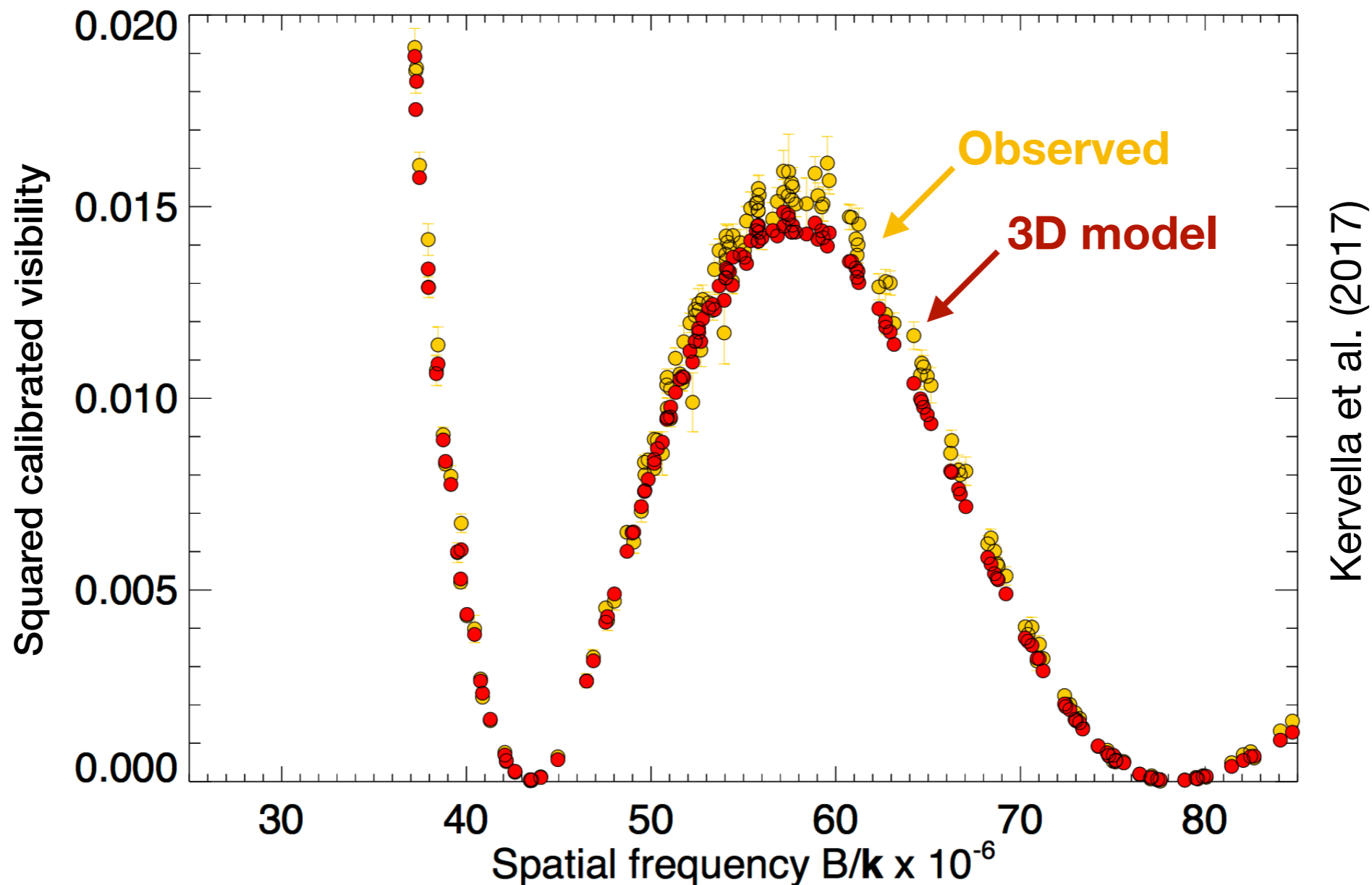
# Interferometric tests

- $\alpha$  Cen B (K0 dwarf), VLT/PIONIER observations in H band ( $1.65 \mu\text{m}$ ) (Kervella, Bigot, Gallene, & Thévenin 2017)
- Observed squared visibilities: two side lobes



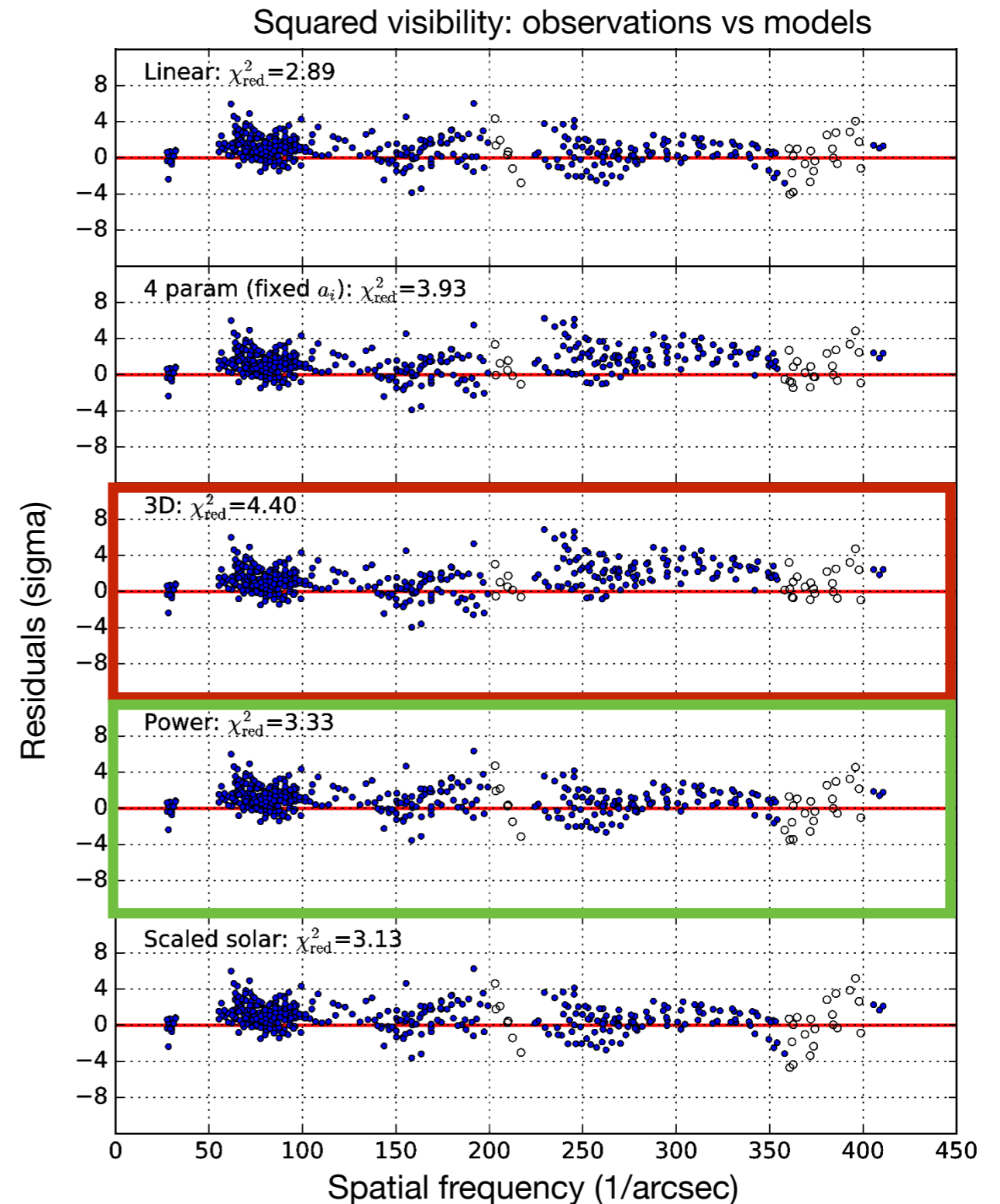
# Interferometric tests

- $\alpha$  Cen B, limb darkening from 3D (and 1D) models: **predicted interferometric square visibility of first side lobe lower than observed**



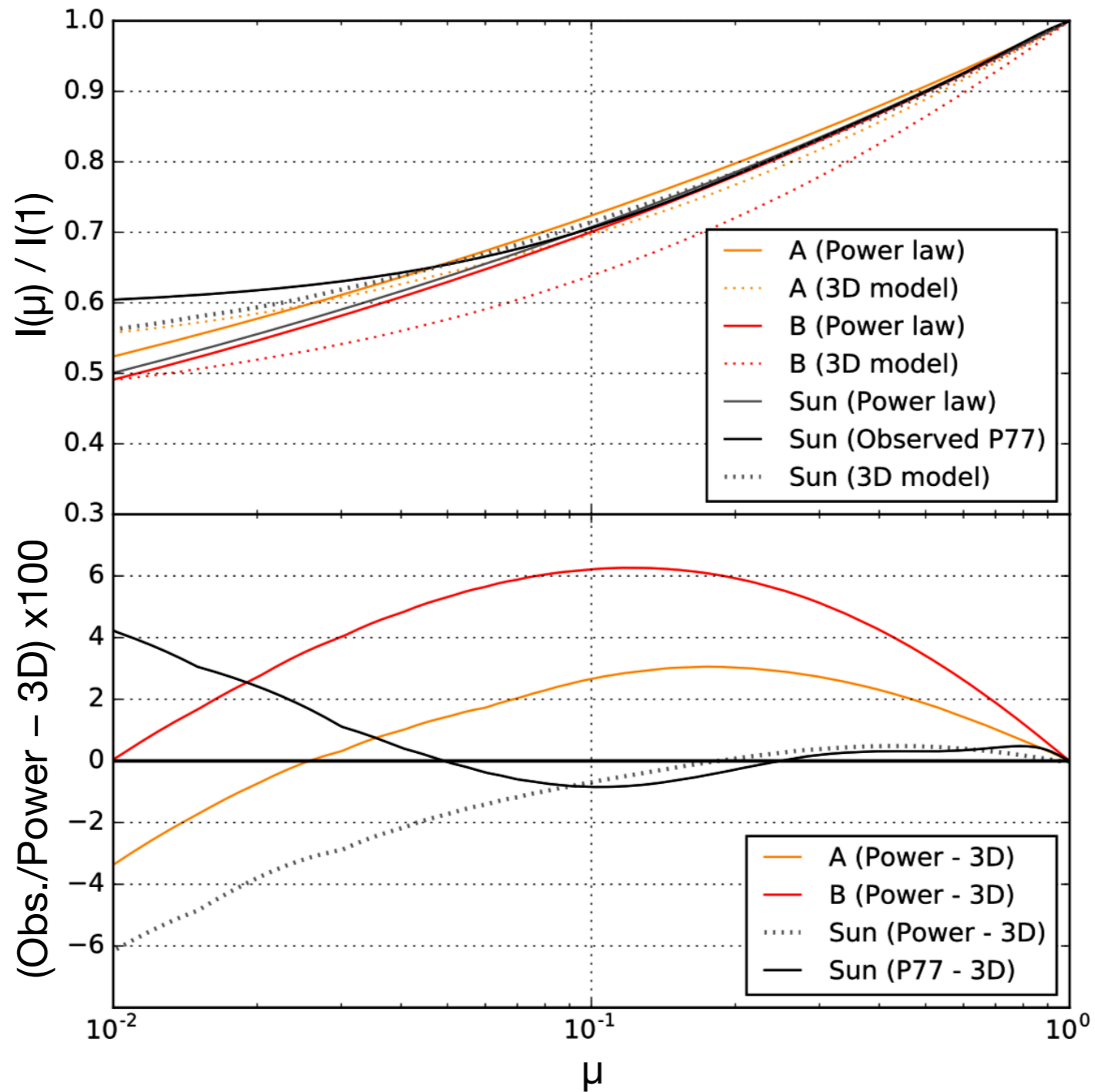
# Interferometric tests

- $\alpha$  Cen B, comparative study of limb darkening models (Kervella et al. 2017)
- Simple power-law model ( $\mu^a$ ) for limb darkening performs better than models



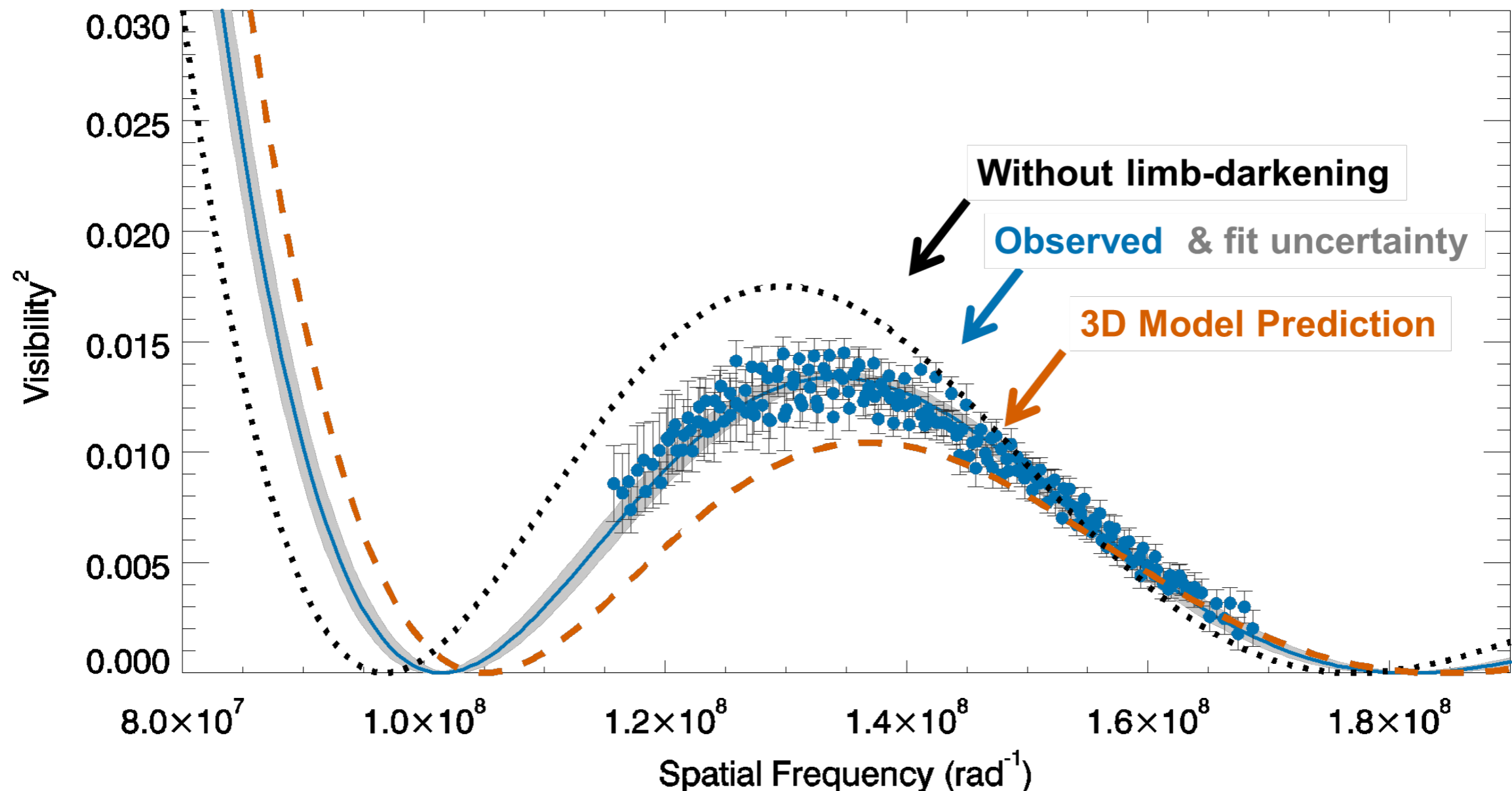
# Limb-darkening models

Comparison between  $\alpha$  Cen A & B and the Sun, power law vs 3D (Kervella et al. 2017)



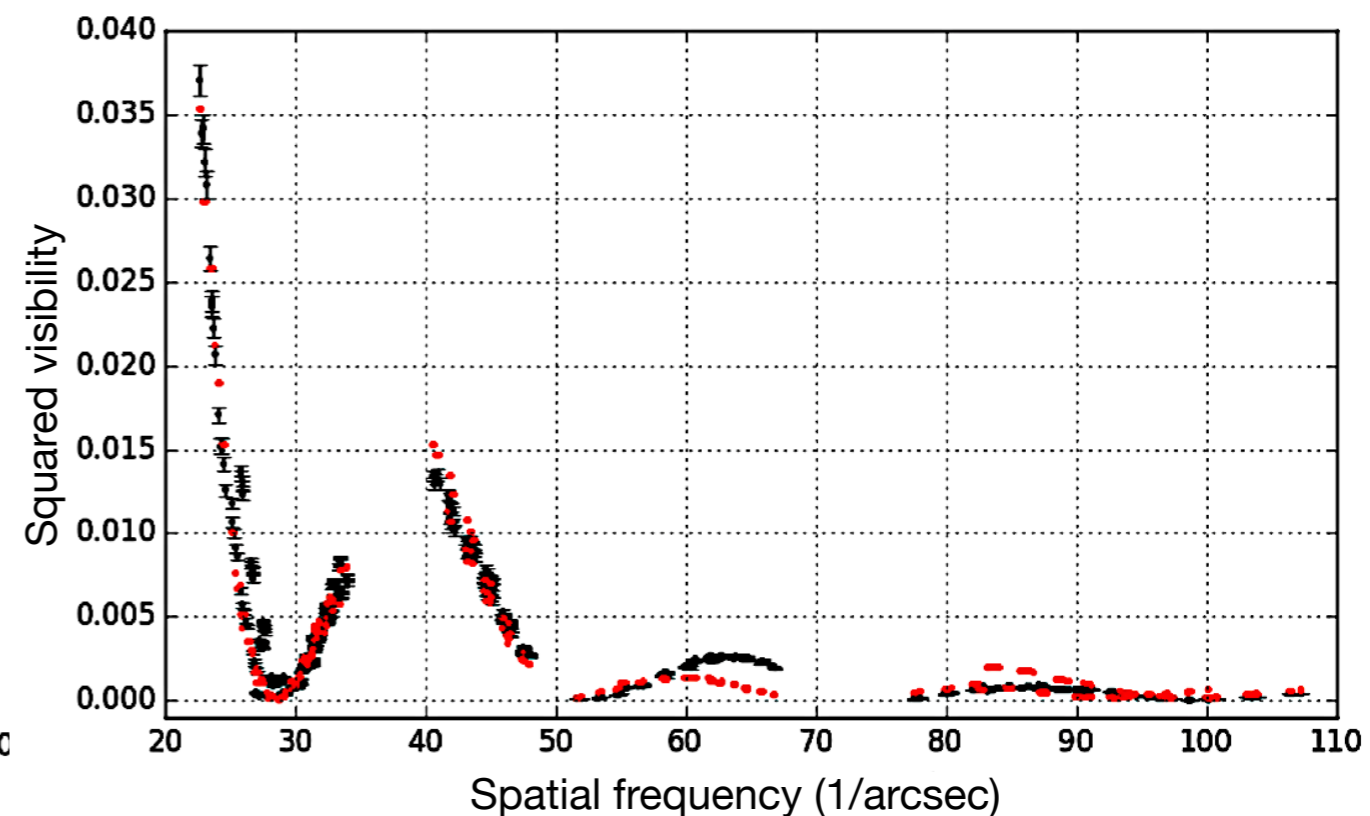
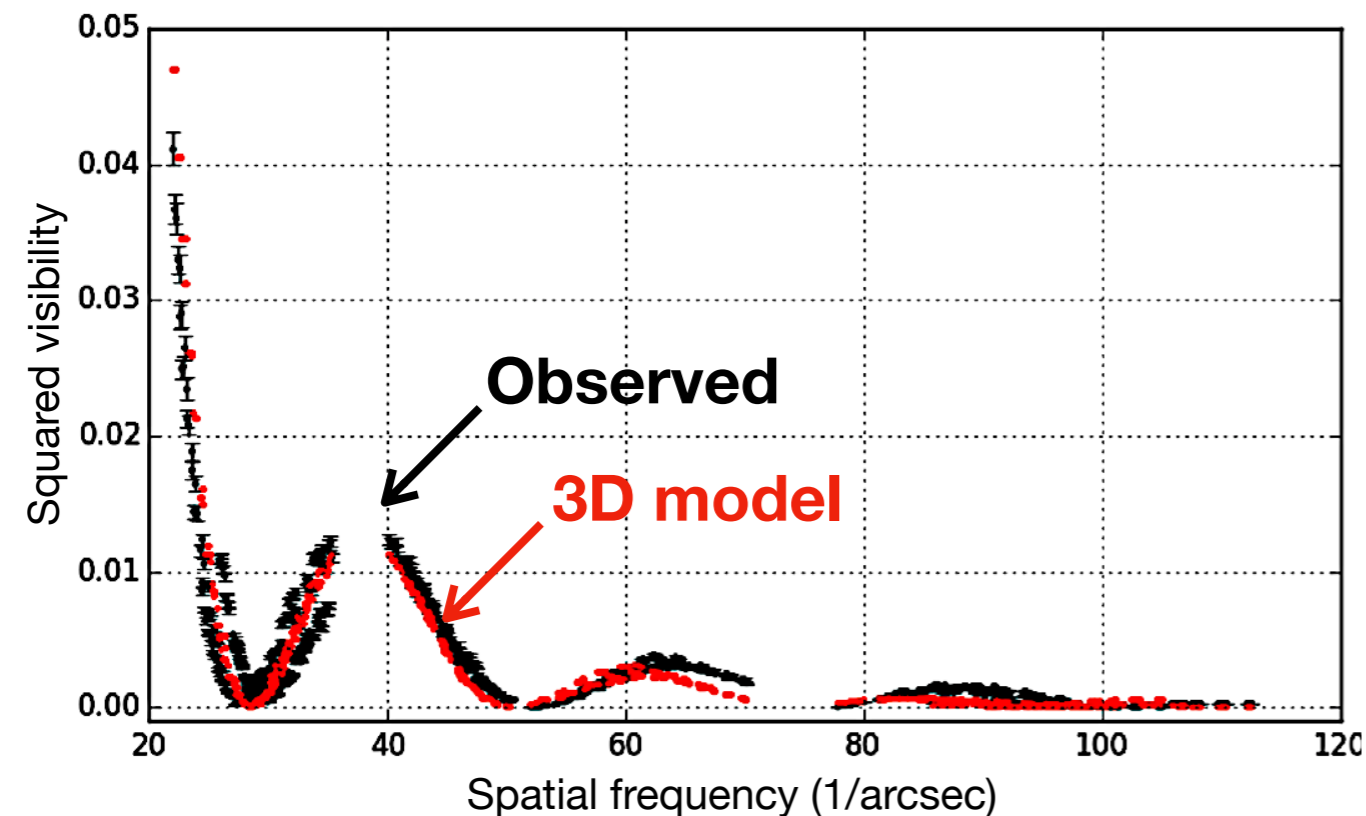
# Interferometric tests (2)

- Evolved star  $\eta$  Cep (K0 subgiant), CHARA/PAVO observations, visible wavelengths (T. White et al., in prep.)



# Interferometric tests (2)

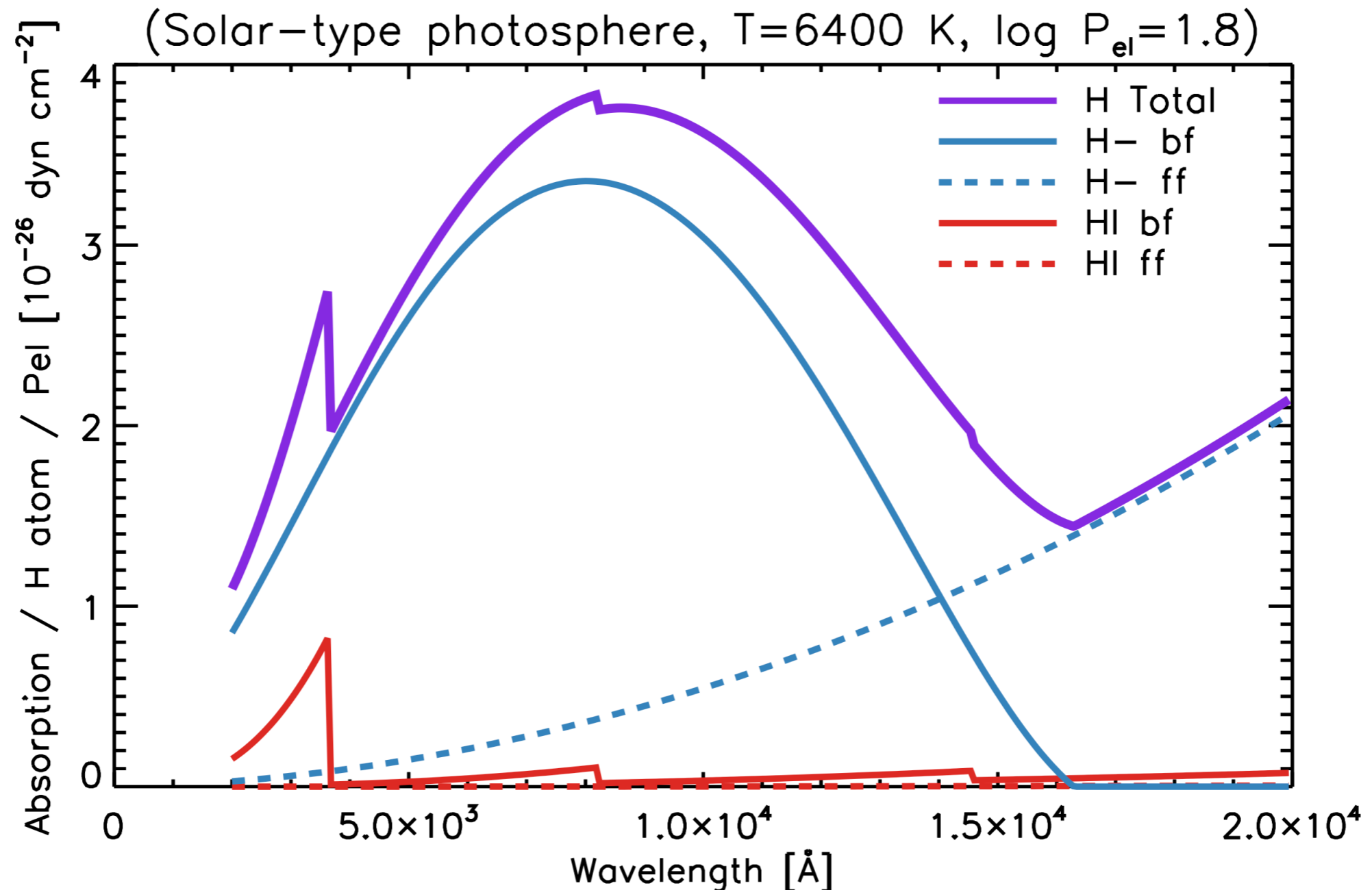
- Antares (M0.5 supergiant), VLT/PIONIER observations (Montargès et al. 2017)
- 3D “star-in-the-box” CO<sup>5</sup>BOLD simulations (Freytag et al. 2012)
- Squared visibilities and angular stellar diameter determinations: dependence on epoch and sampled position angles





# Continuous opacities

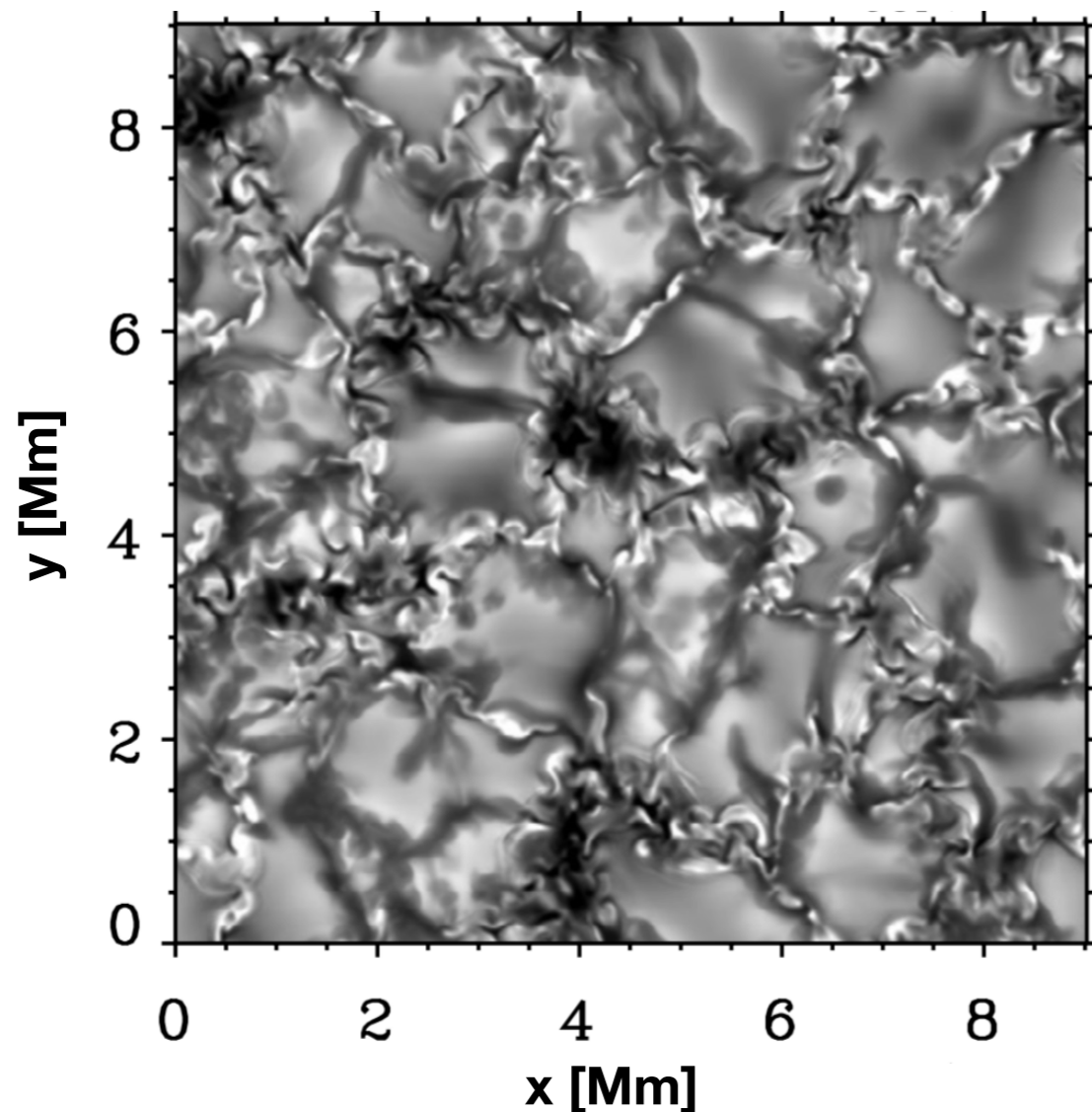
- H- bound-free: cross-section? (direct/indirect) effects of departures from LTE on H- or electron number density?



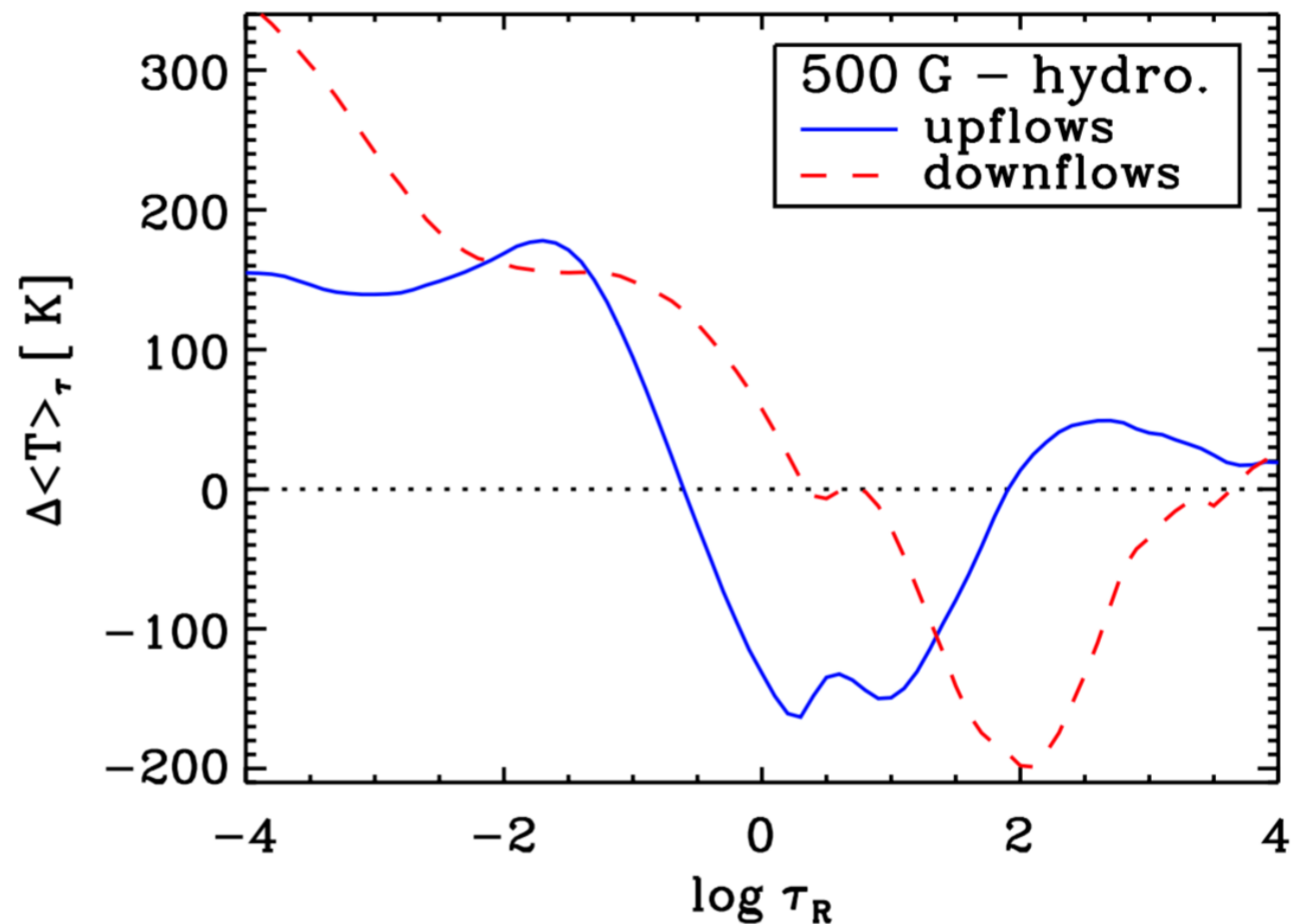
# Magnetic fields

- Effect of vertical magnetic fields on physical stratification in main-sequence stars (e.g. Beeck et al. 2015; Rempel 2014; Pereira et al. 2013)

G2V star,  $B_0=500\text{G}$

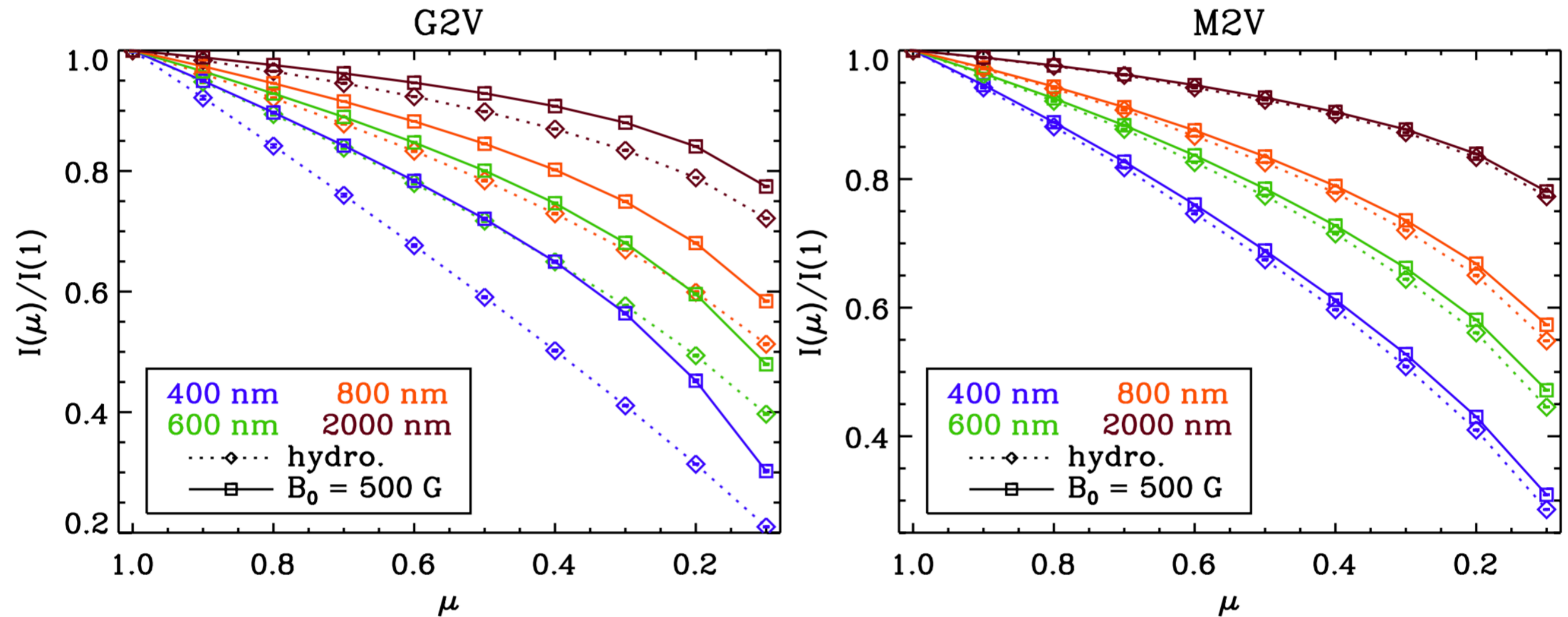


G2V



# Magnetic fields

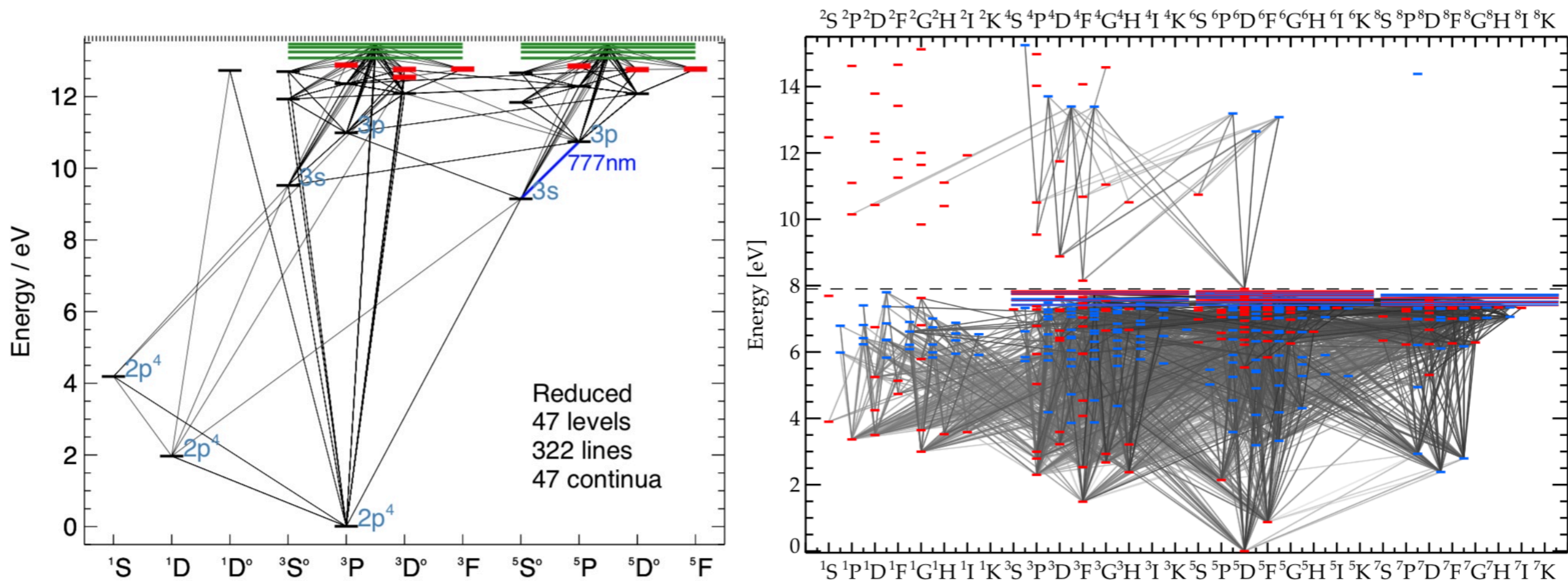
Effect of magnetic fields on limb-darkening curves



Beeck et al. (2015)

# Spectral lines, CLV and non-LTE

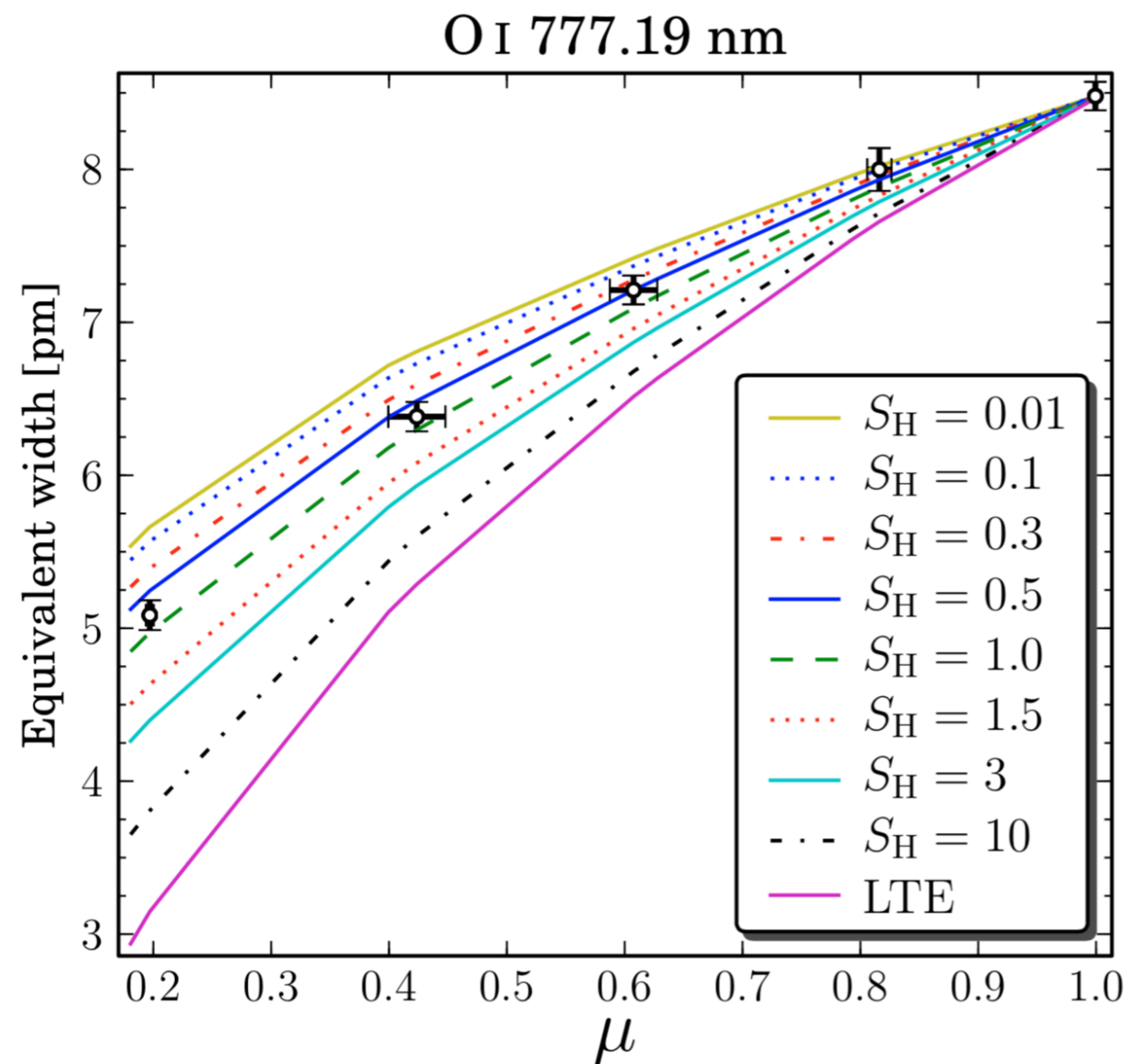
- Departure of atomic level populations and radiation field from local thermodynamic equilibrium (LTE) in stellar atmospheres
- BALDER/MULTI3D code (Amarsi et al. 2016, 2018), scaling to large model atoms (C, O, Si, Fe, ...)



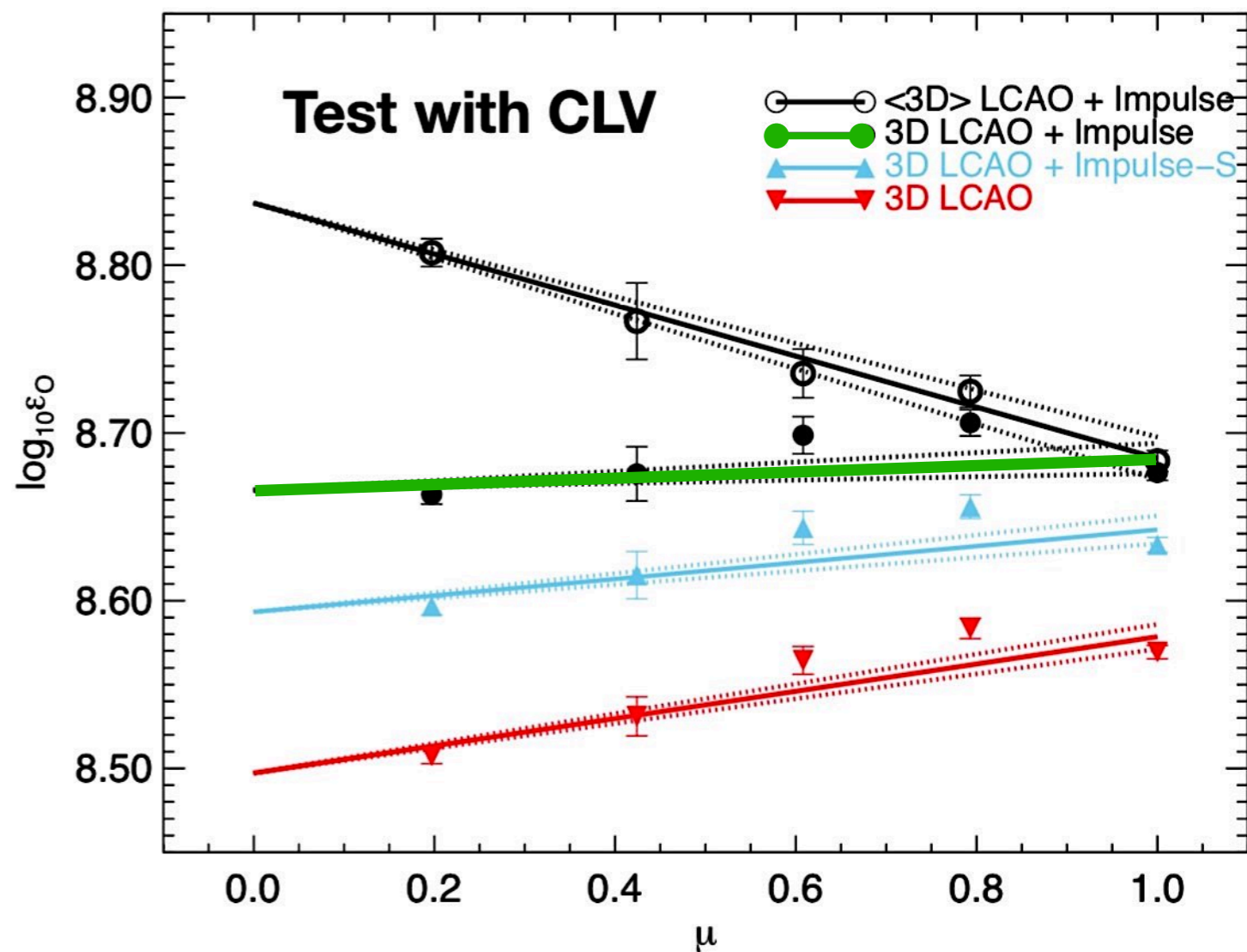
Model O and Fe atoms (Amarsi et al. 2018; Lind et al. 2017)

# Oxygen lines and CLV (Sun)

- 3D non-LTE O abundance and CLV: sensitivity to model atmosphere structure and atomic/electronic collisions



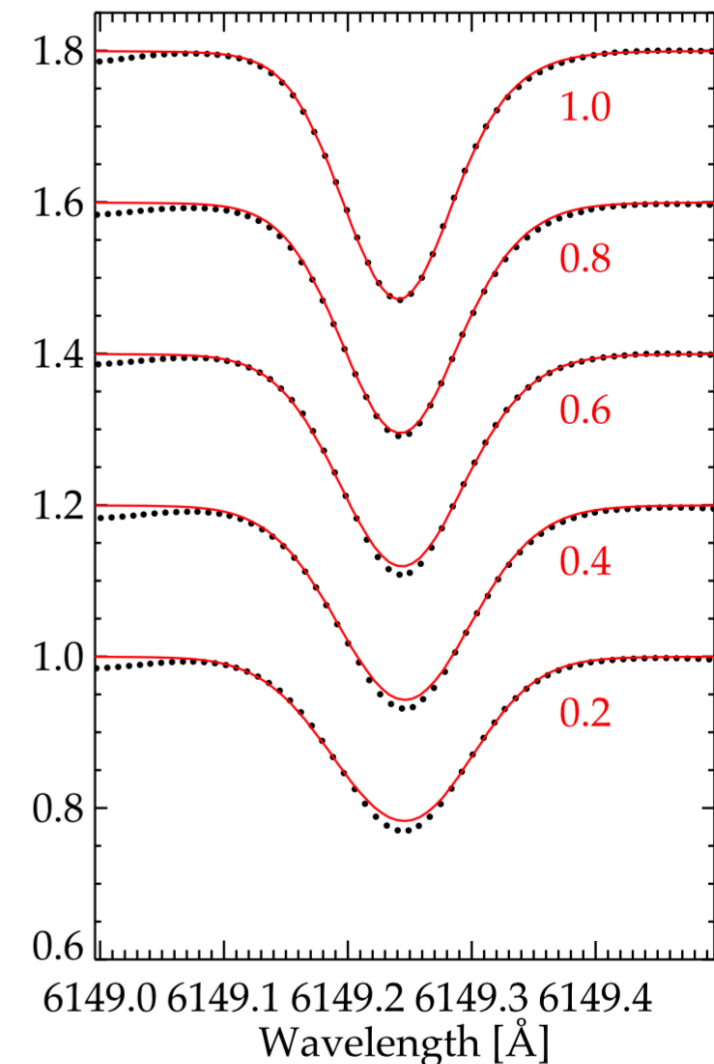
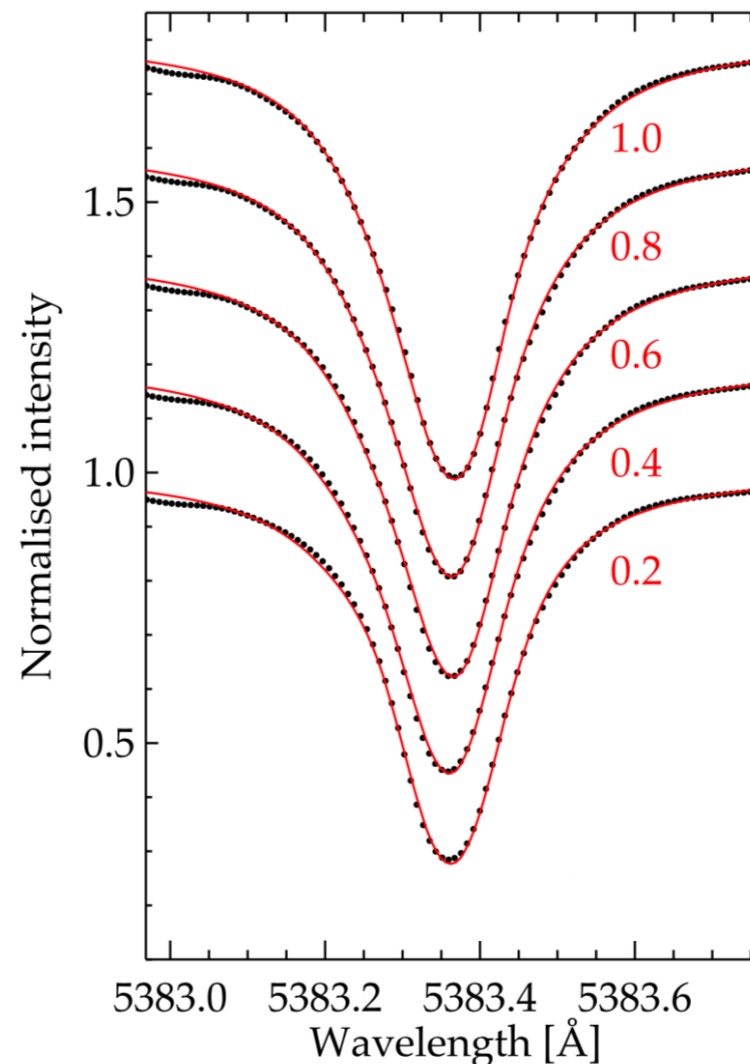
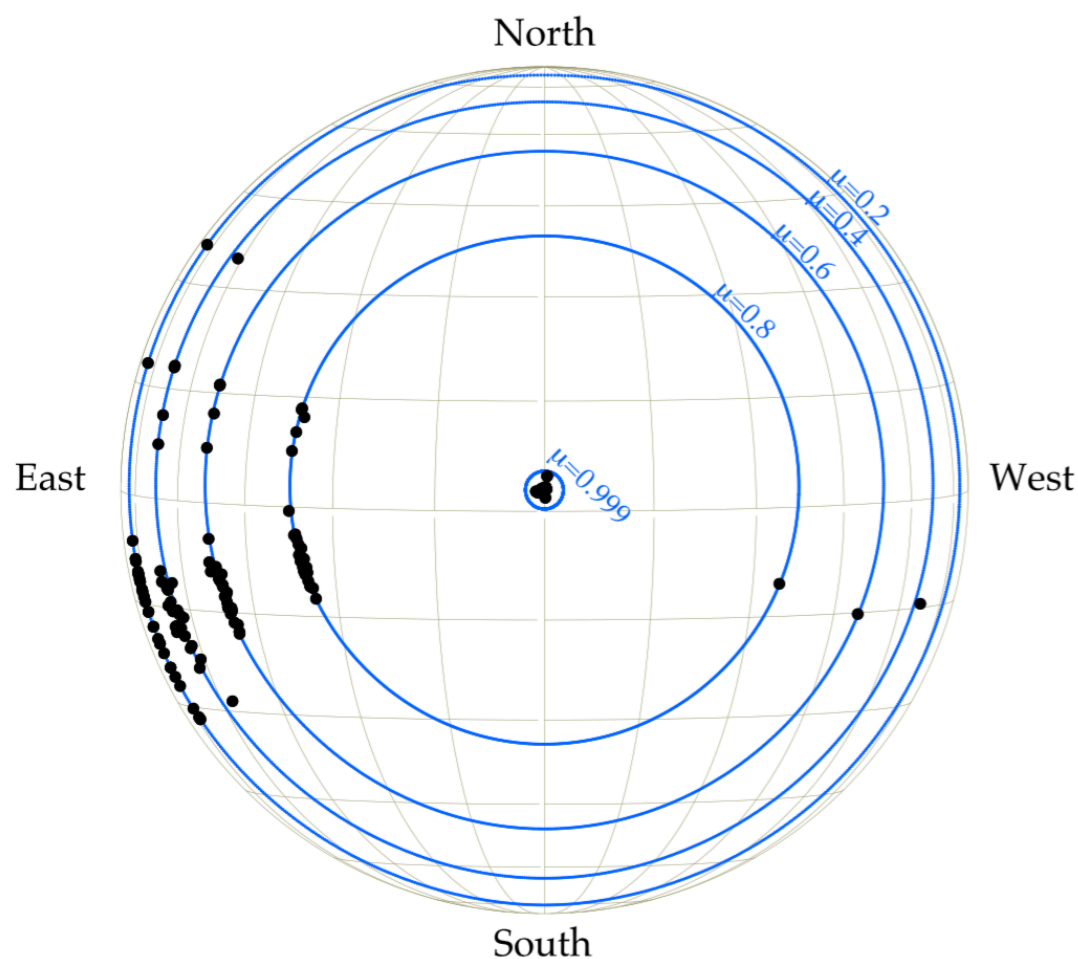
Pereira et al. (2013)



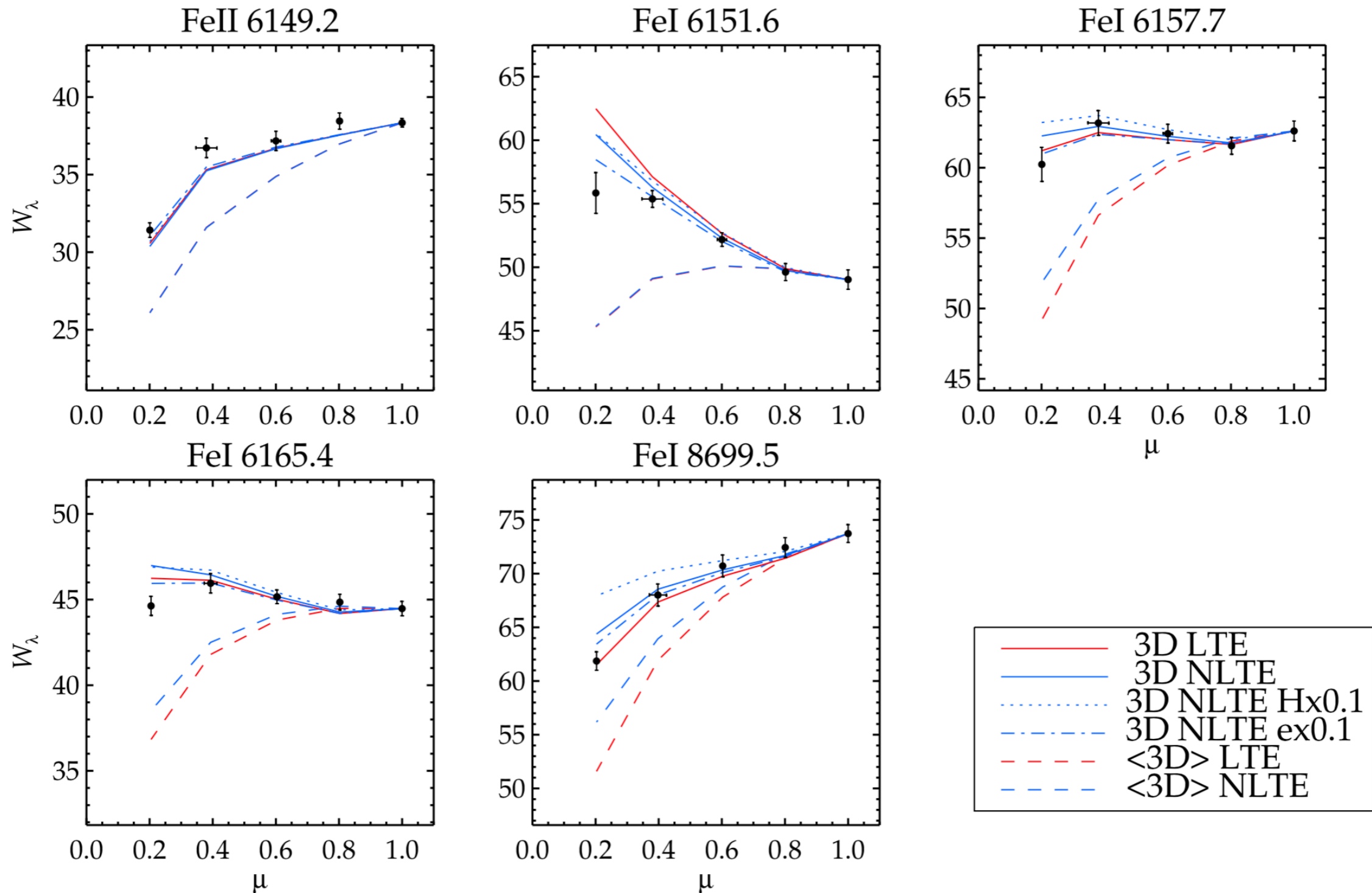
Amarsi et al. (2018)

# Iron lines and CLV (Sun)

- Observations with Swedish Solar Telescope (Lind et al. 2017)
- 3D non-LTE modelling of Fe I and II spectral lines
- Quantum mechanical Fe+H collisional rates

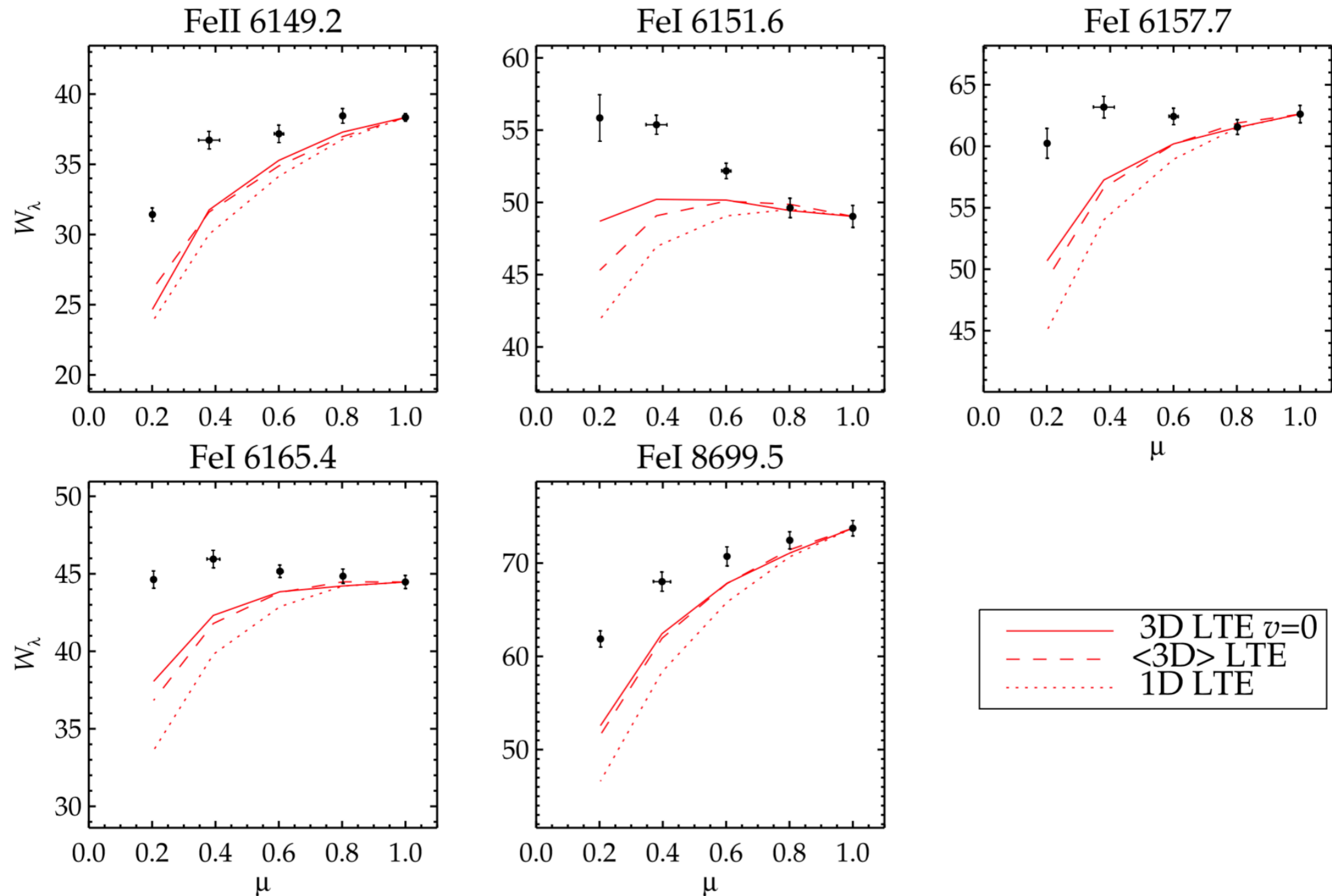


# Iron lines and CLV (Sun)



Lind et al. (2017)

# Iron lines and CLV (Sun)



Lind et al. (2017)



# Collaborators

## **WP122200 and STAGGER teams:**

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